

Atmosphere Considerations and Advantages

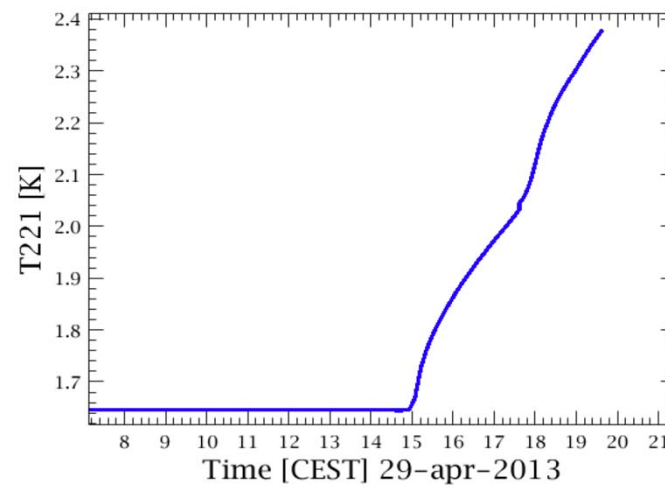
Steve Lord

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(Ah, Goodbye, Herschel....)



Scope of this talk

- I have modeled the atmosphere in the FIR using a program I wrote at NASA-Ames – ATRAN.
- ATRAN is about one of 4 such programs available
- For a comparison see Guan, X, Stutzki, J et al 2012
- They do differ by several percent, but this is not relevant for the purposes of this talk.

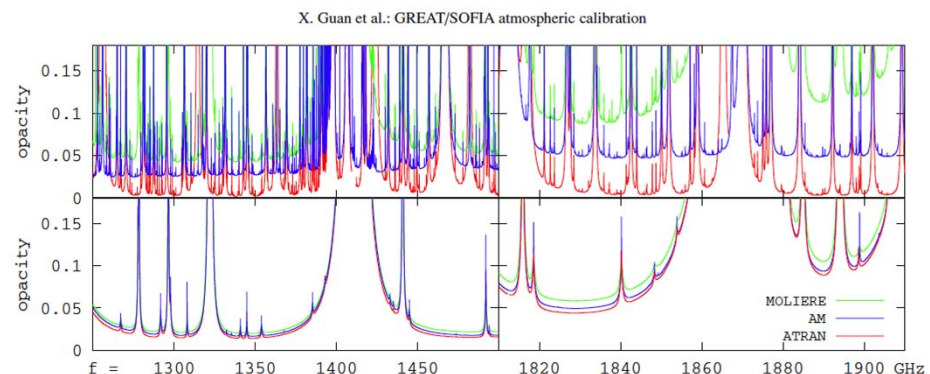
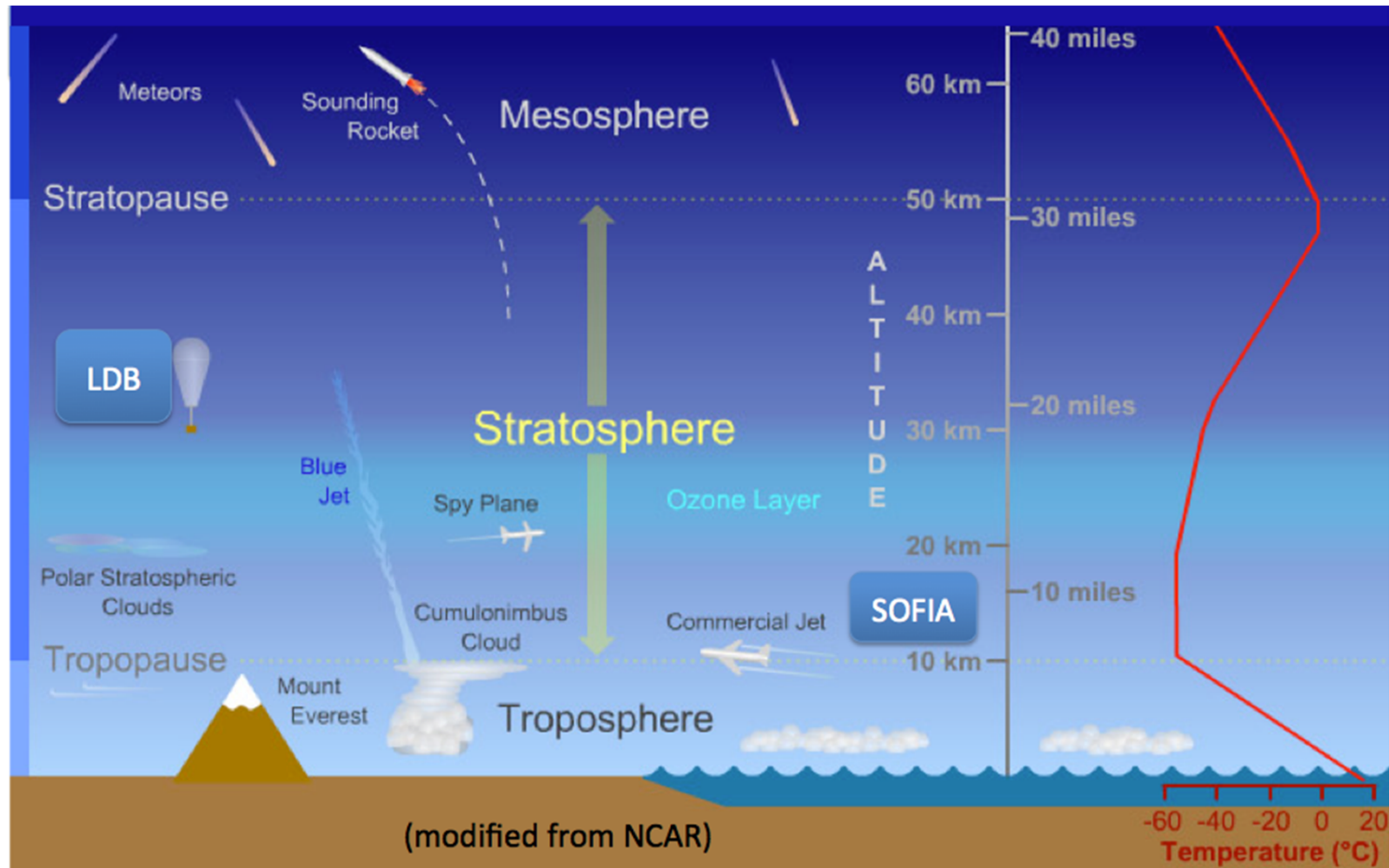


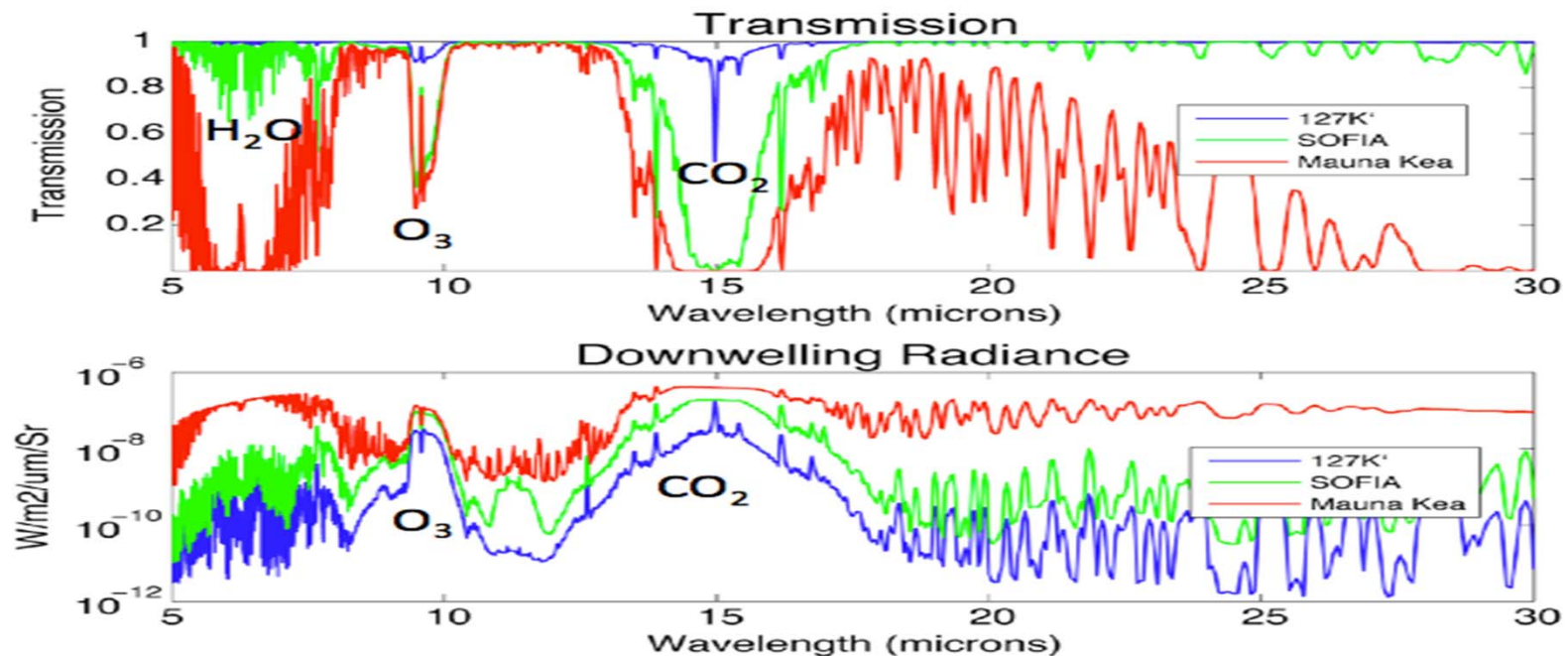
Fig. 1. Comparison of the dry (*top*) and wet (*bottom*) opacity coefficients as defined in Eqs. (6), (A.29) and (A.27) for the frequency range of GREAT channels L1 (*left*) and L2 (*right*), calculated from the atmospheric models AM, ATRAN and MOLIERE (see text), showing the opacity for a pressure altitude of $p_{\text{amb}} = 133$ Torr and a $p_{\text{uv}} = 10 \mu\text{m}$. The largest discrepancy is in the quasi-continuous dry-atmospheric opacity due to collision-induced absorption (CIA), which is completely neglected in ATRAN and differs by about a factor of 2 between MOLIERE and AM.

The regime: 20,000' to 95,000'

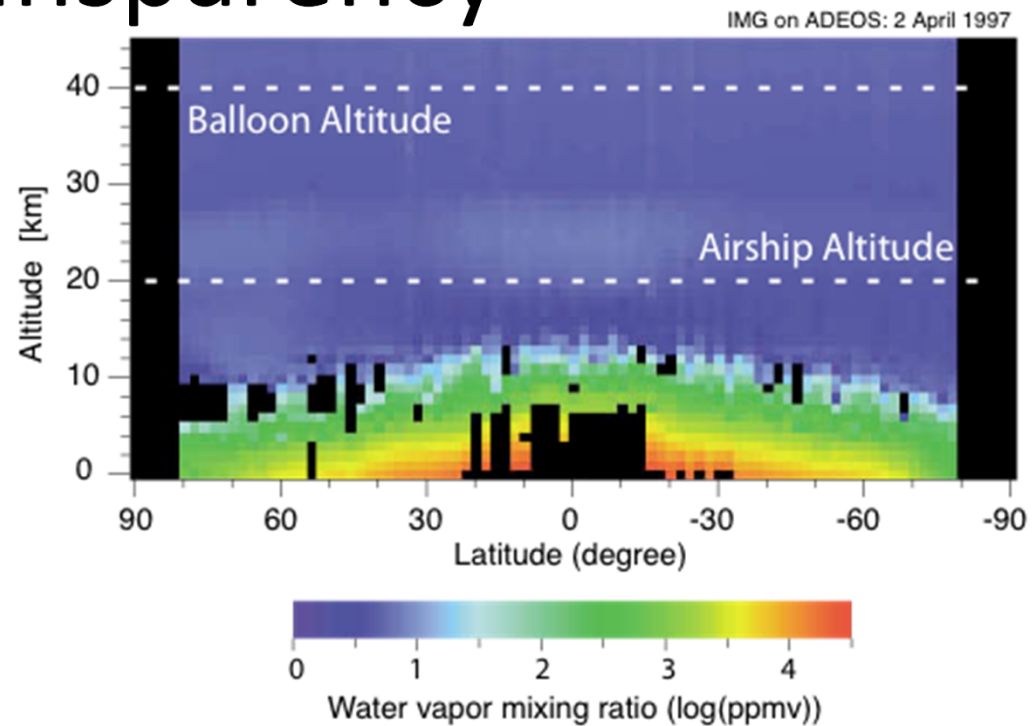
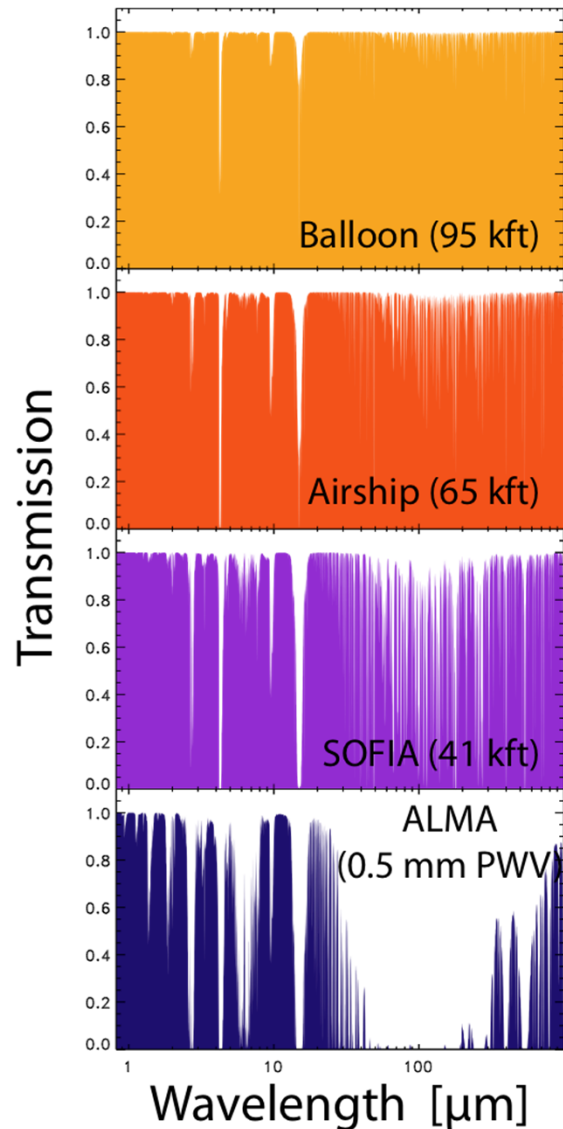


Scope (Continued)

- I will show models (that I have used) for astronomical investigations, and apply them at various altitudes. I only consider emission below: (c/o E. Young)
- Relevant to BLIP NEP of detectors



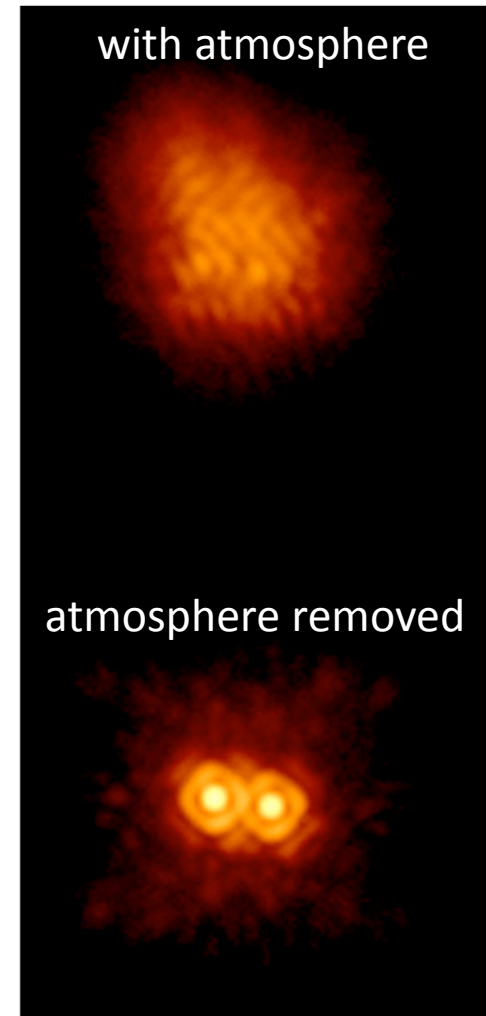
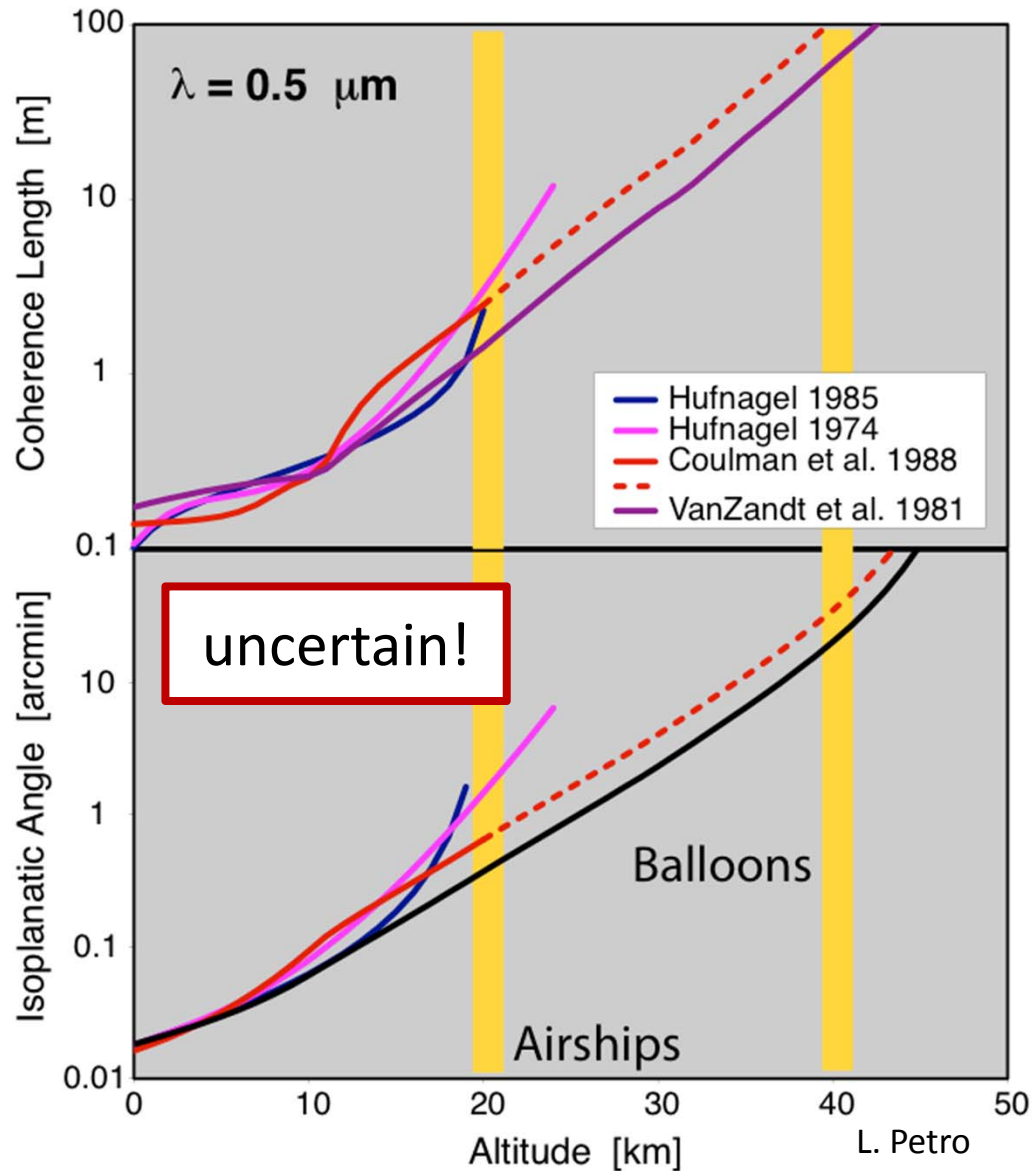
Astronomy: Atmospheric Transparency



Steven Lord, Caltech
(custom made)

balloons and airships are
above most water

Astronomy: Image Quality



IW Tau: Beichman & Tanner

Modeling the Transmittance for 3 Observatories:

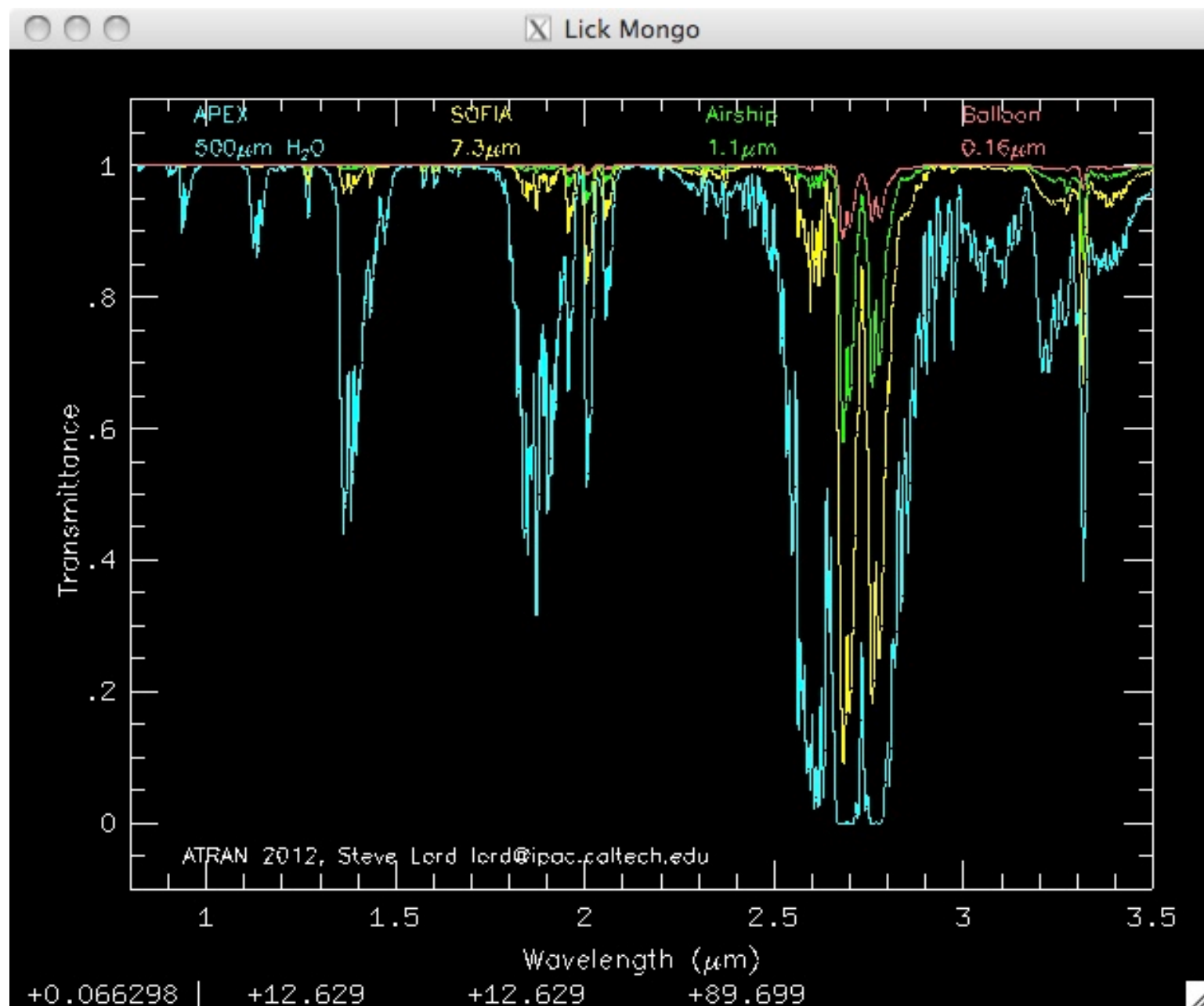
Obs	alt (ft)	H2O um	O3 (cm2)	ZA

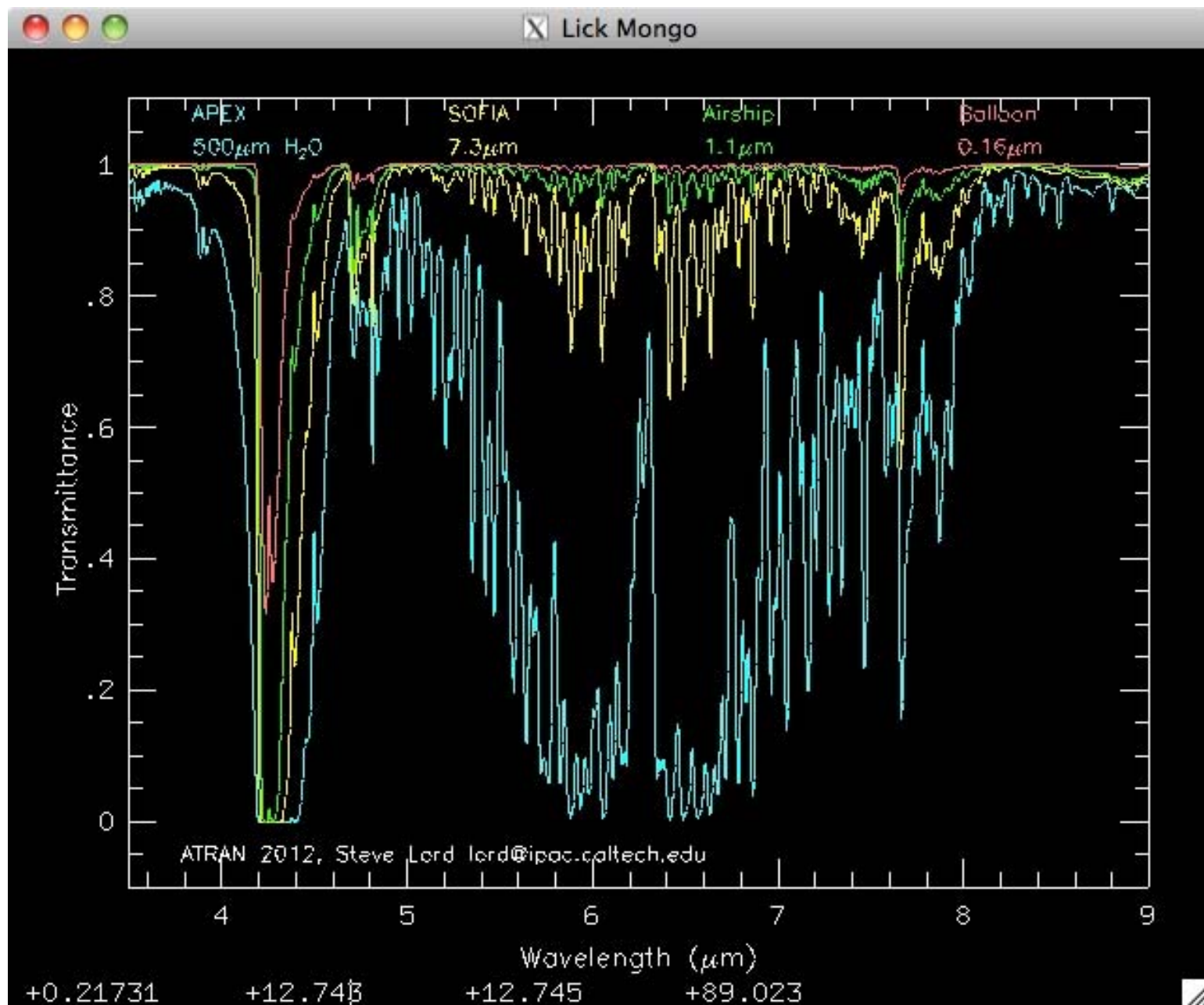
Alma	16597.	500.000	0.91E+19	0.
SOFIA	41000.	7.300	0.91E+19	0.
Airship	65000.	1.100	0.91E+19	0.
Balloon	95000.	0.157	0.38E+19	0.

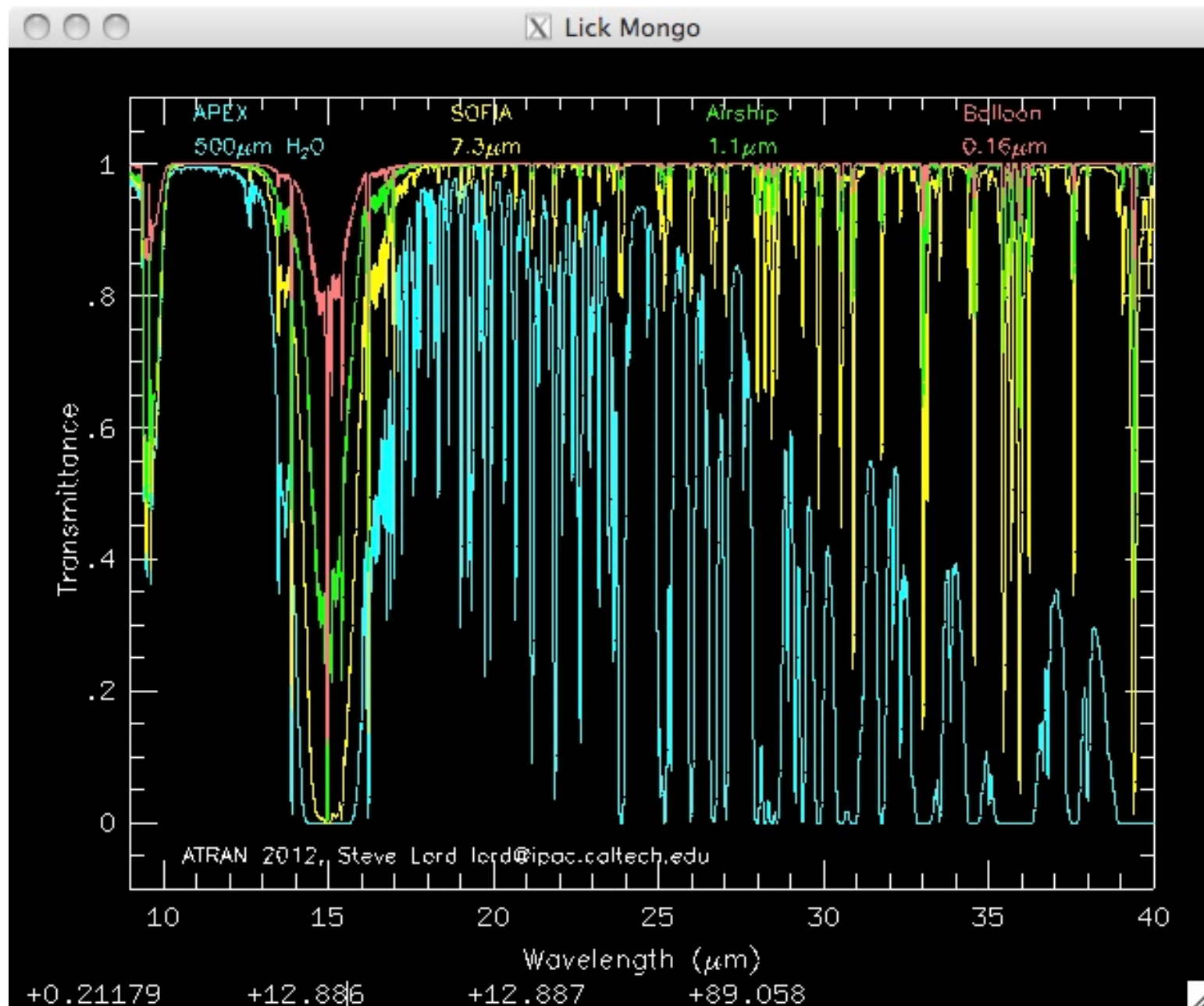
Resolution used for the plots below:

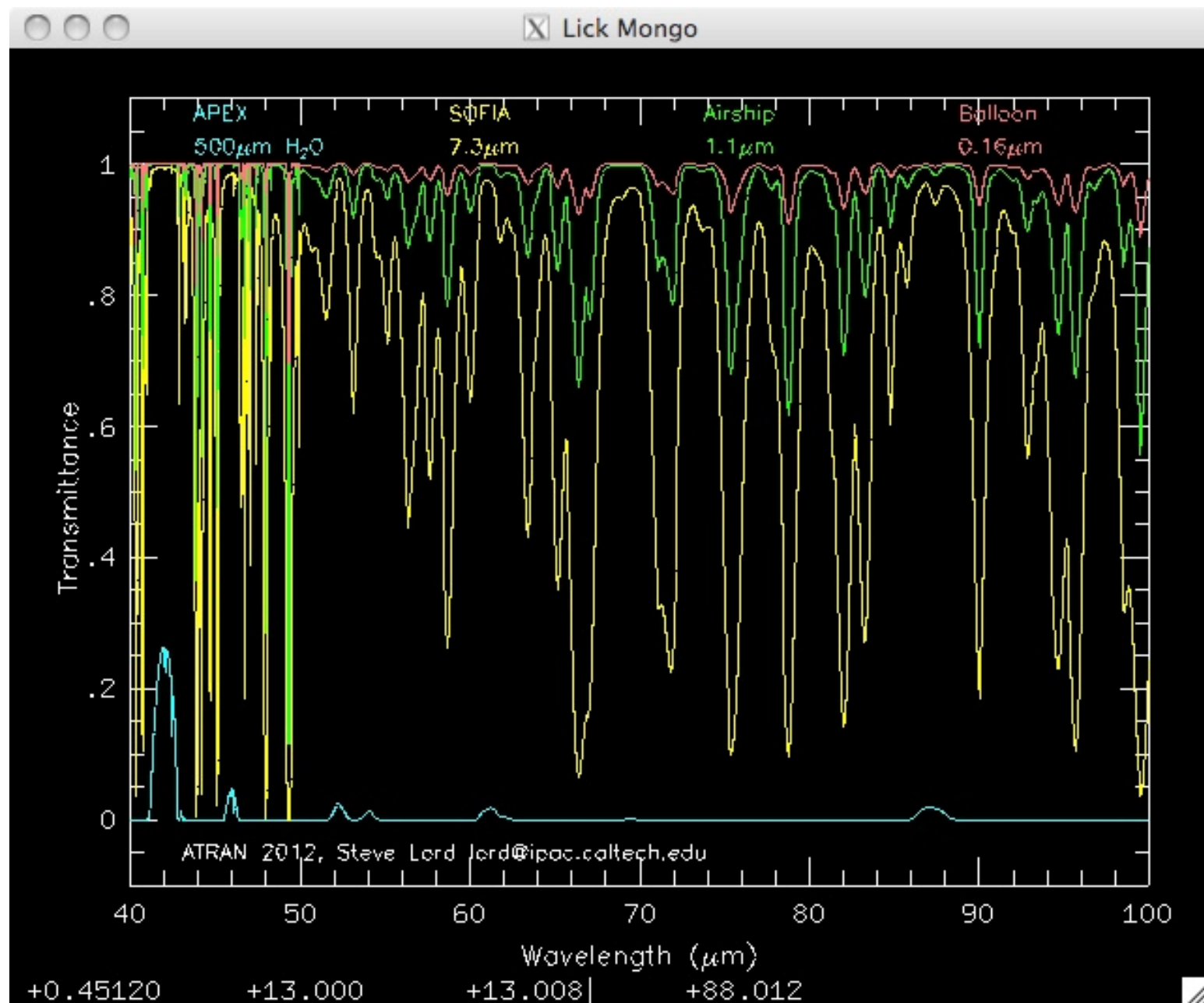
lam 1	lam 2	resolution
(um)	(um)	(um)

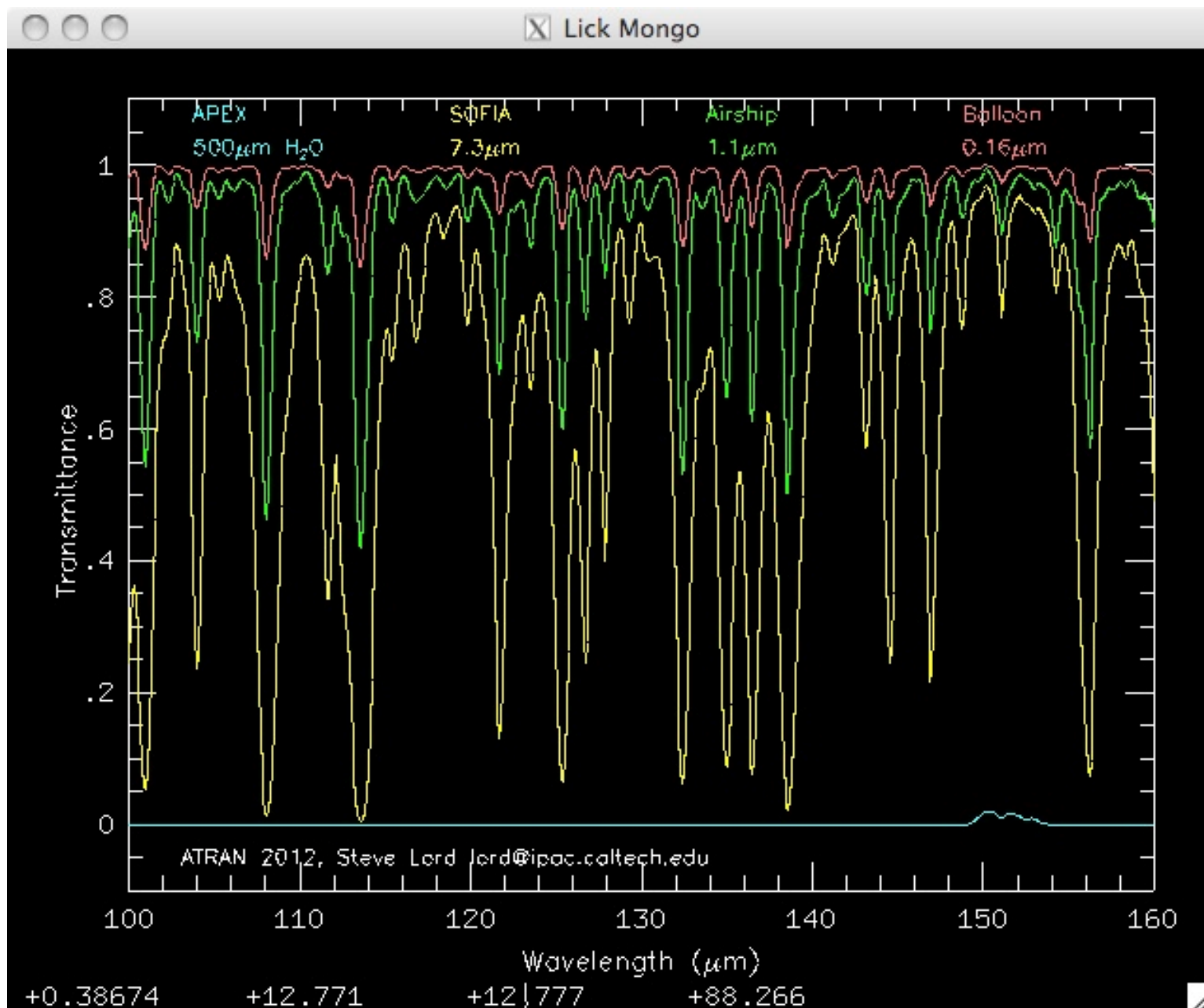
0.82	3.5	0.005
3.5	9	0.01
9	40	0.05
40	100	0.1
100	160	0.5
150	300	0.5
300	1000	1.0

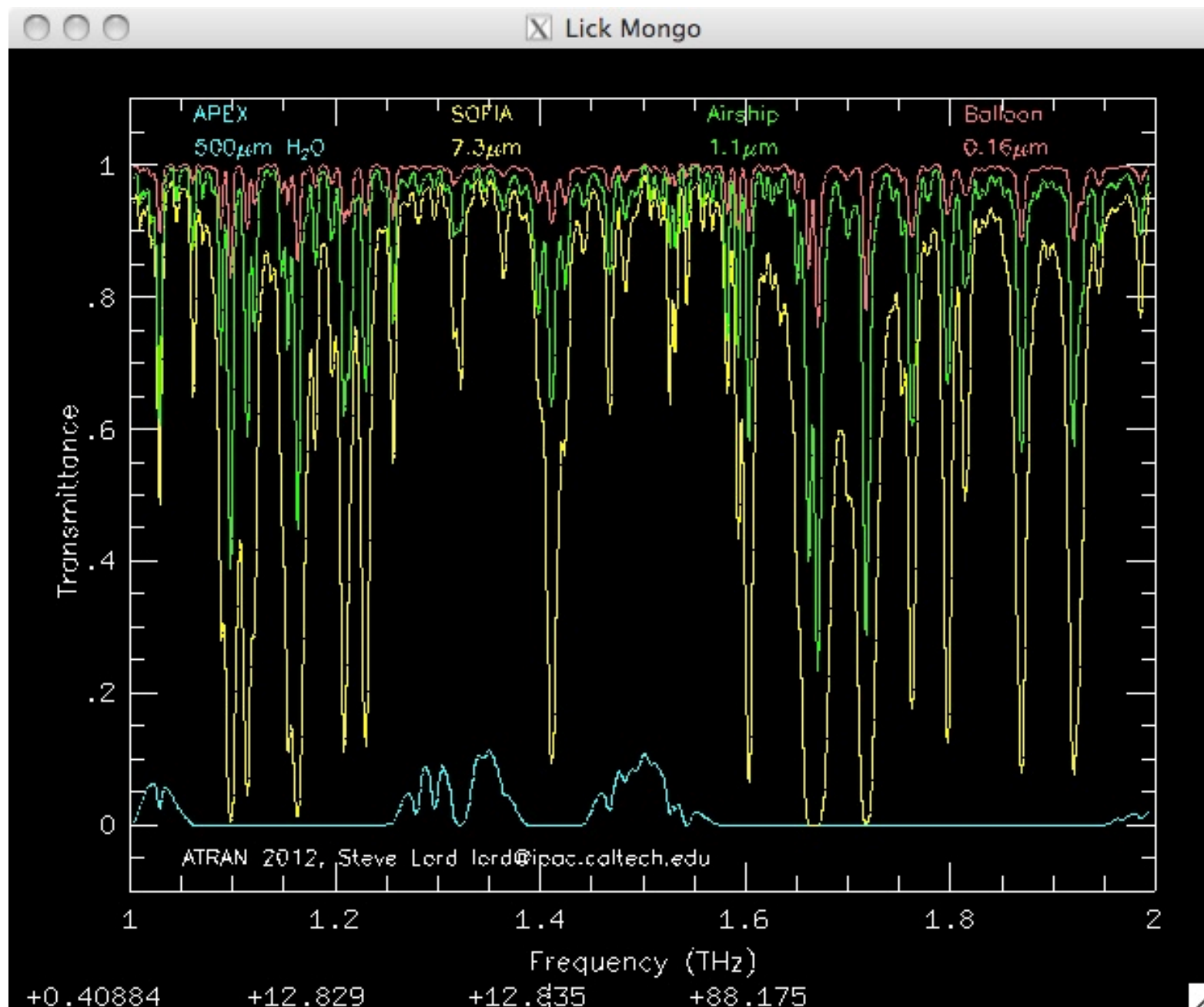


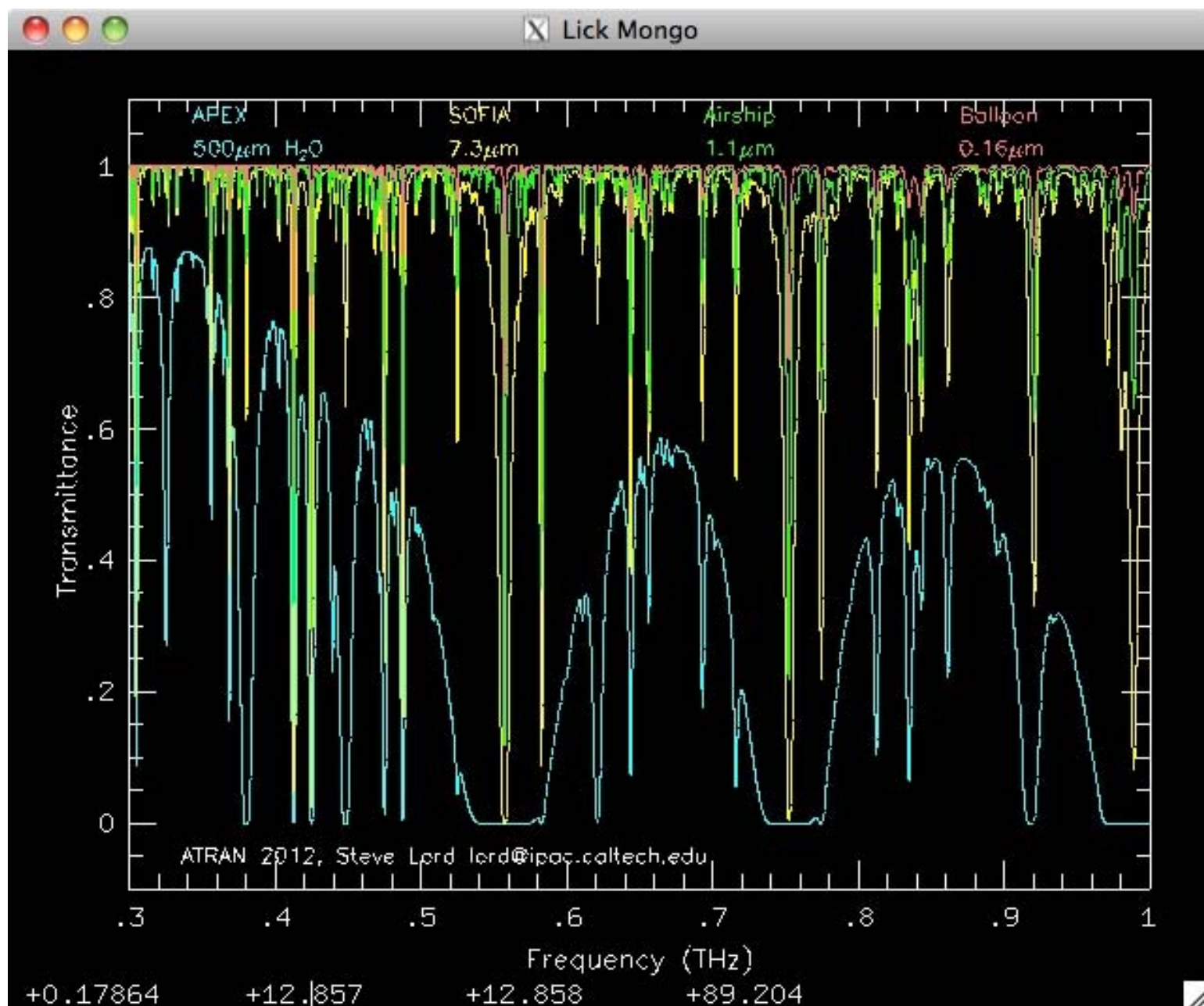












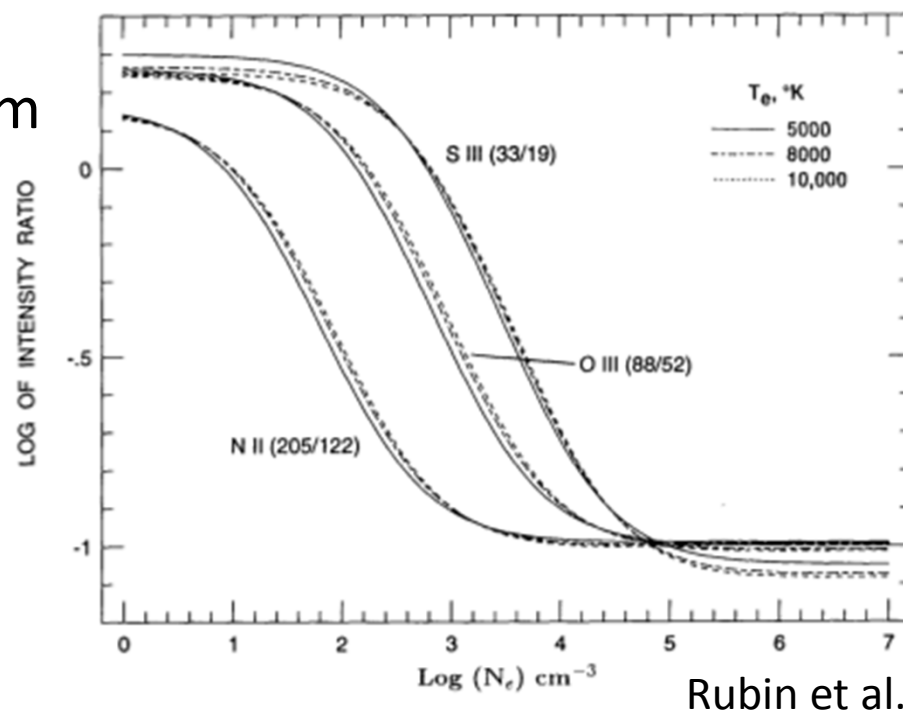
Lines Presented in the next slides:

1. SIII at wavelength 18.71300 HII region
2. FeIII at wavelength 22.93000 HII region
3. OIV at wavelength 25.87000 AGN
4. FeII at wavelength 25.98820 HII region
5. SIII at wavelength 33.48000 HII region
6. SiII at wavelength 34.81400 HII region, PDR
7. NeIII at wavelength 36.01000 HII region
8. OIII at wavelength 51.81500 HII region
9. NIII at wavelength 57.33000 HII region
10. OI at wavelength 63.18372 PDR
11. OIII at wavelength 88.35600 HII region
12. NII at wavelength 121.89760 Diffuse ISM
13. OI at wavelength 145.52548 PDR
14. CII at wavelength 157.74100 PDR and HII region(+)
15. NII at wavelength 205.17830 Diffuse ISM

(Underlined lines at high redshift too)

Atomic Fine-Structure HII Region Diagnostics

- [SIII] 18 μ m and 33 μ m
- [FeII] 26 μ m and [FeIII] 23 μ m
- [NeIII] 36 μ m
- [OIII] 52 μ m and 88 μ m
- [NIII] 57 μ m

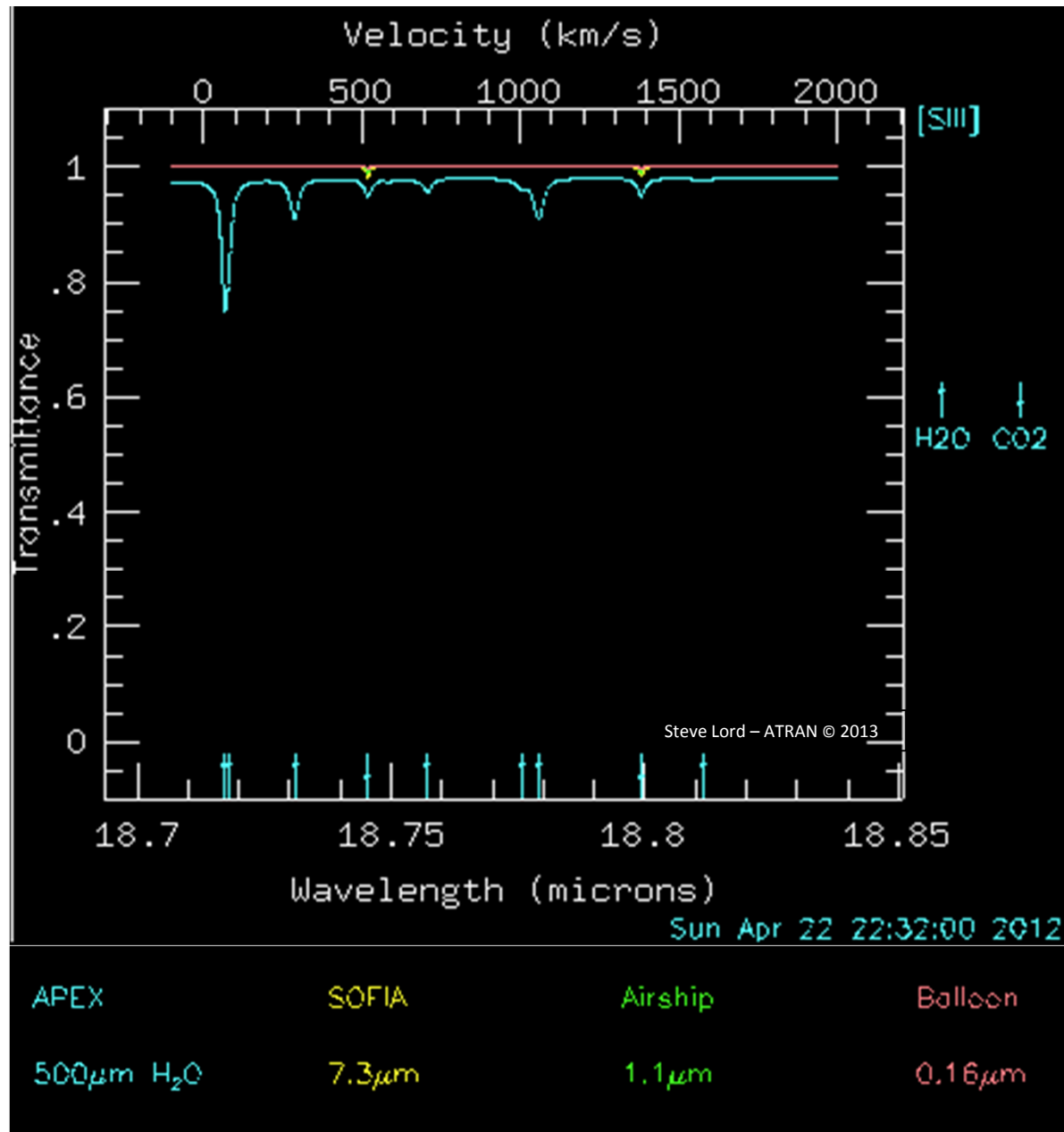


Measure electron density, stellar effective temperature, abundance ratios, abundance gradients, PNe properties,

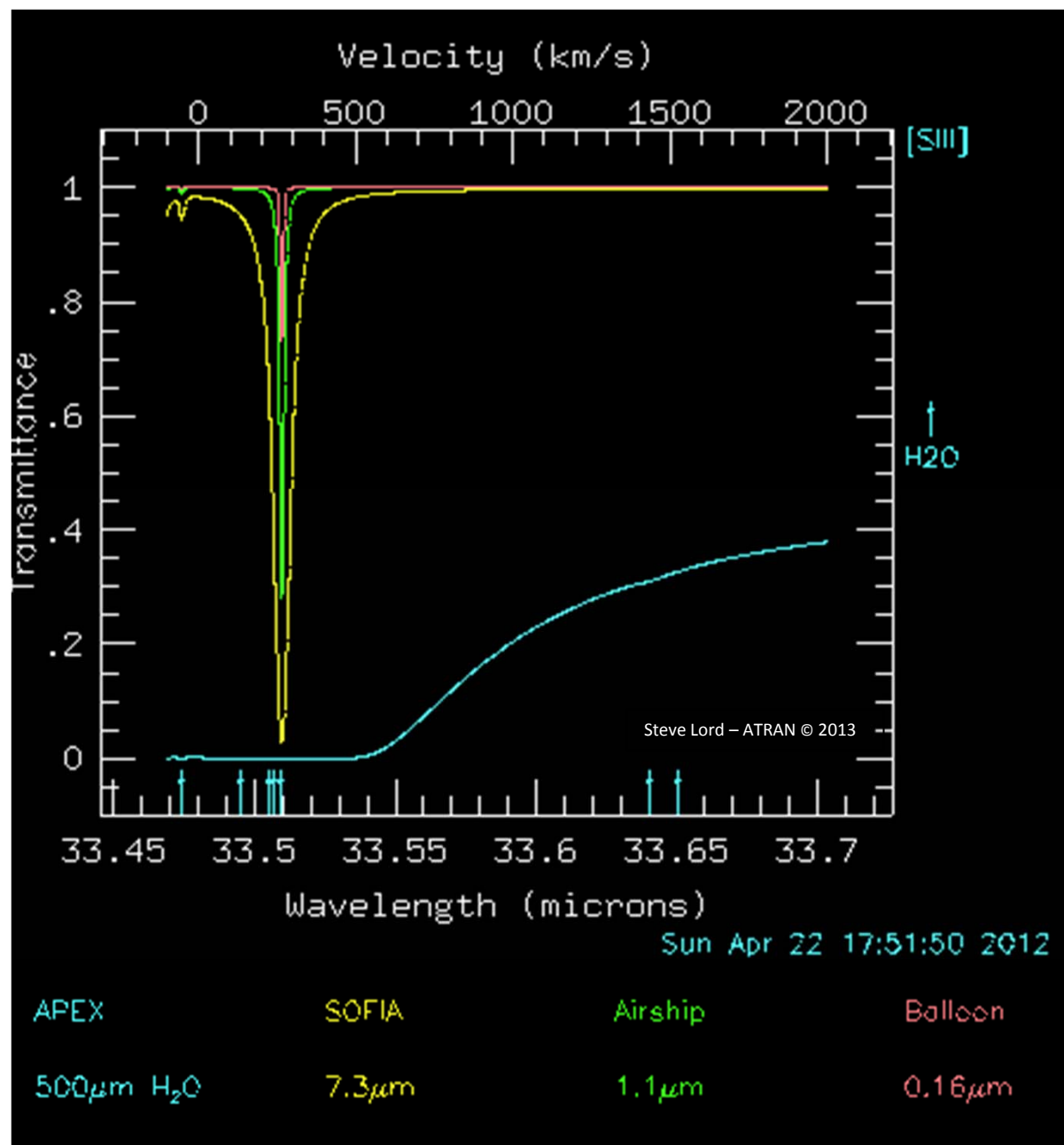
SNRs...

Steve Lord

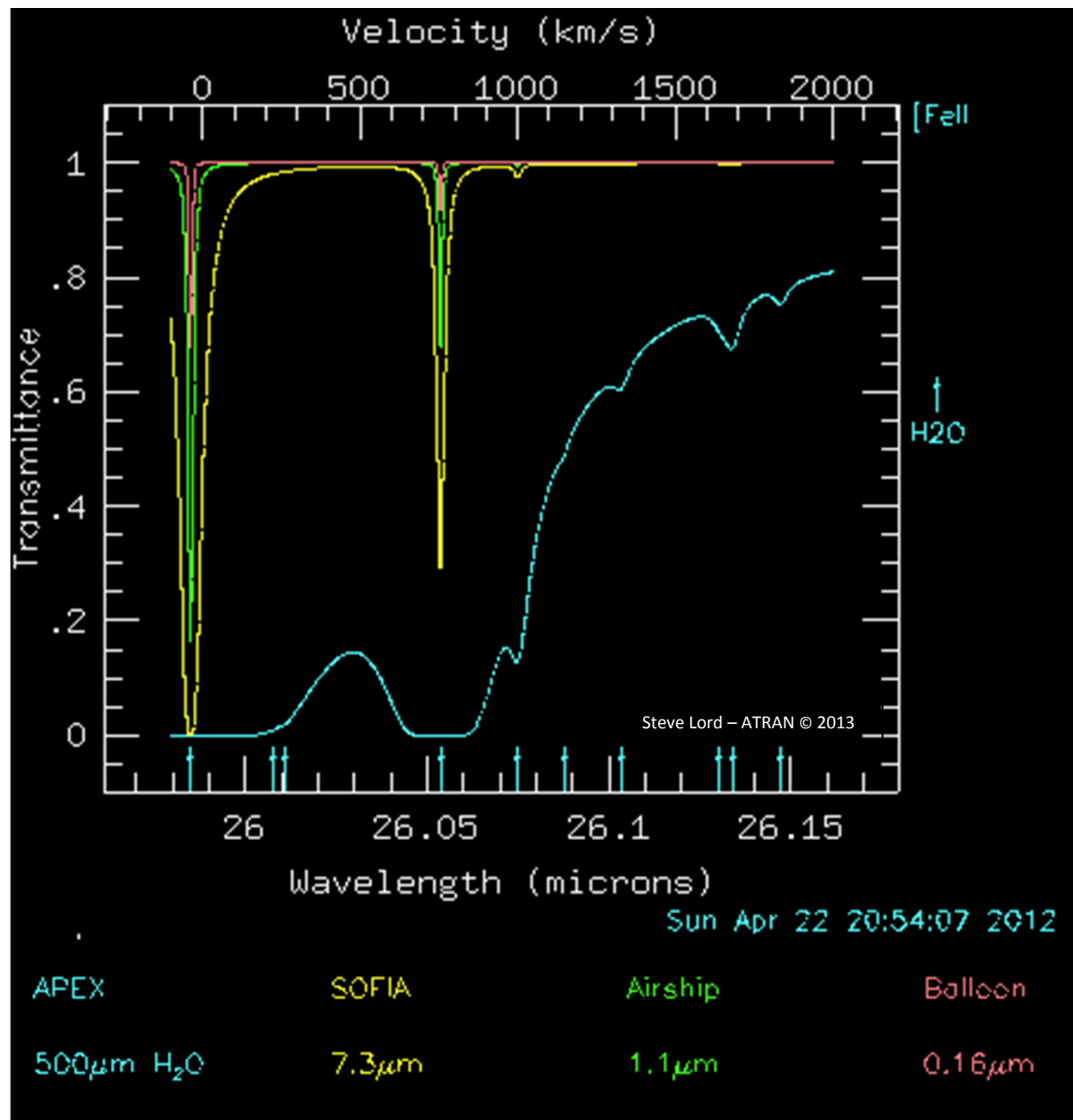
KISS Airship Workshop - Pasadena -
4/30/2013

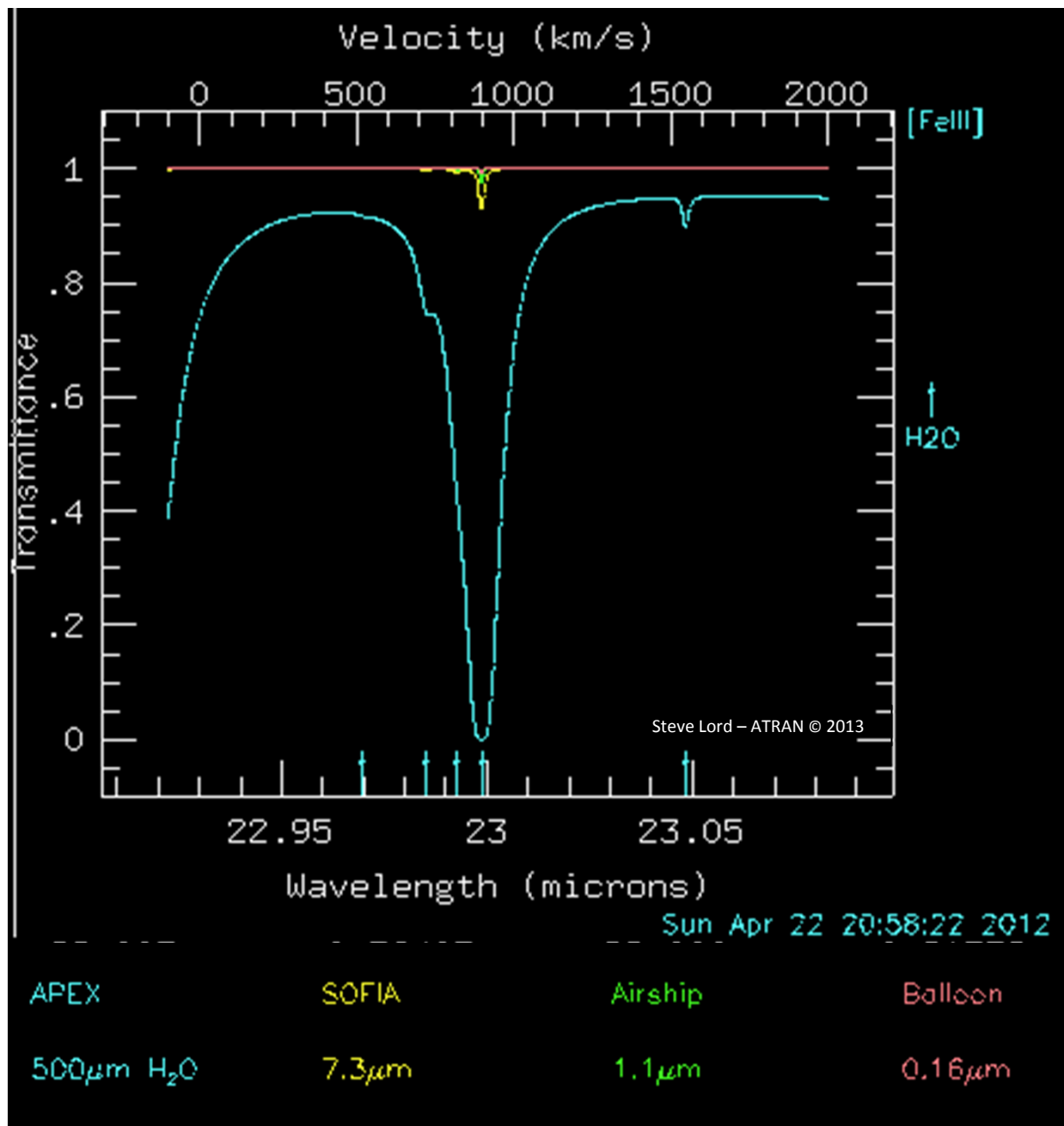


[SIII] 19um

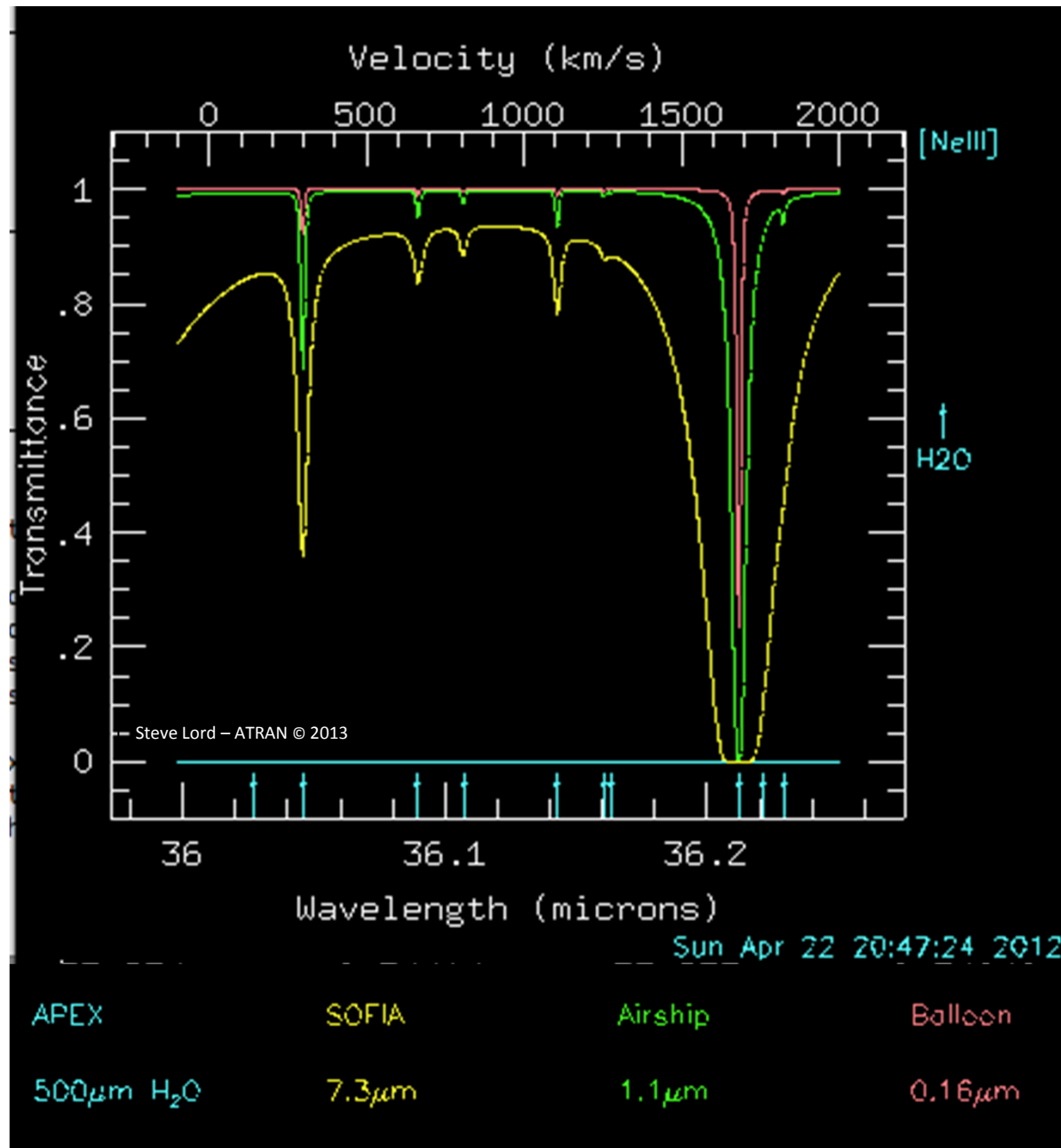


[SIII] 33um



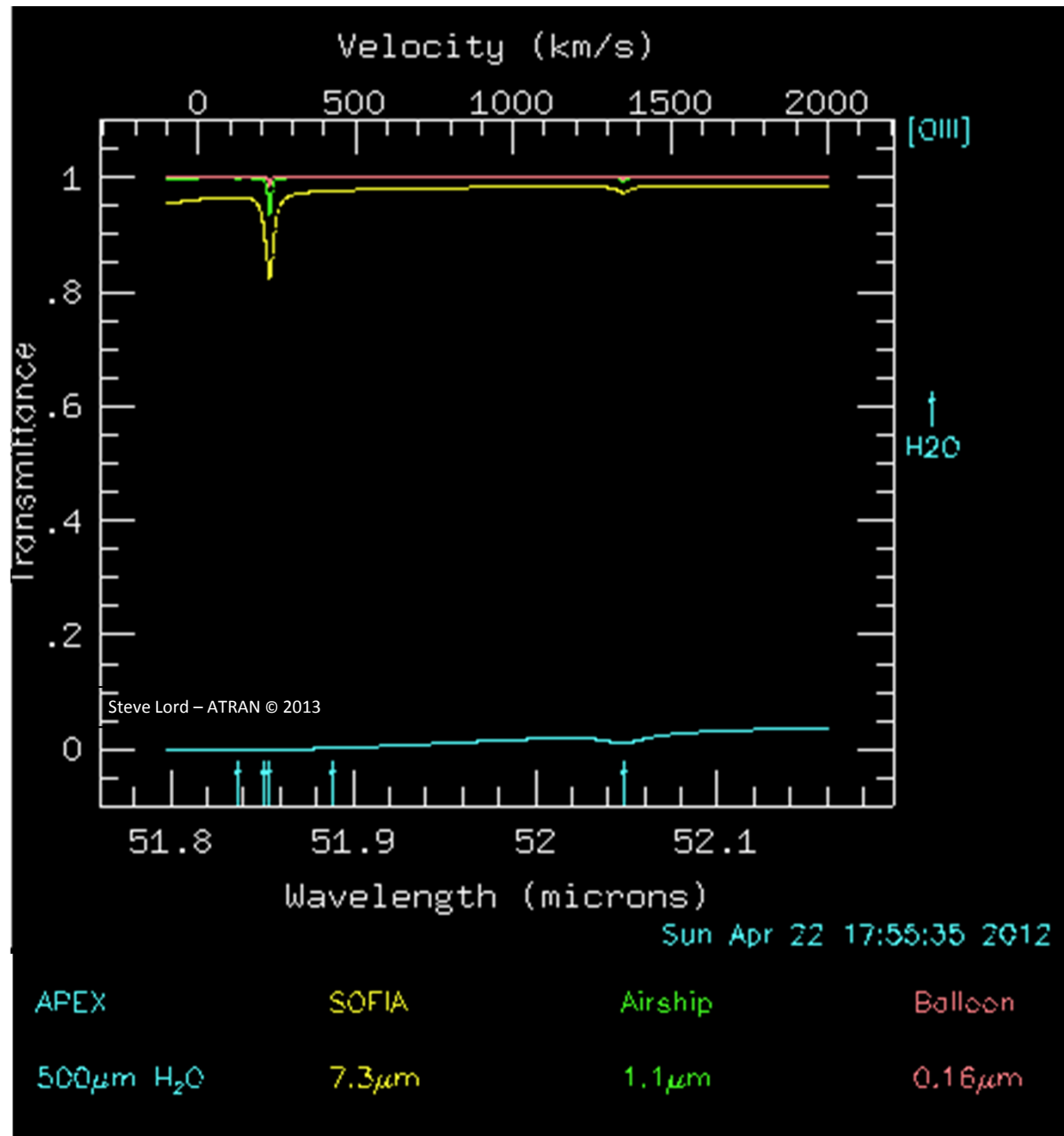


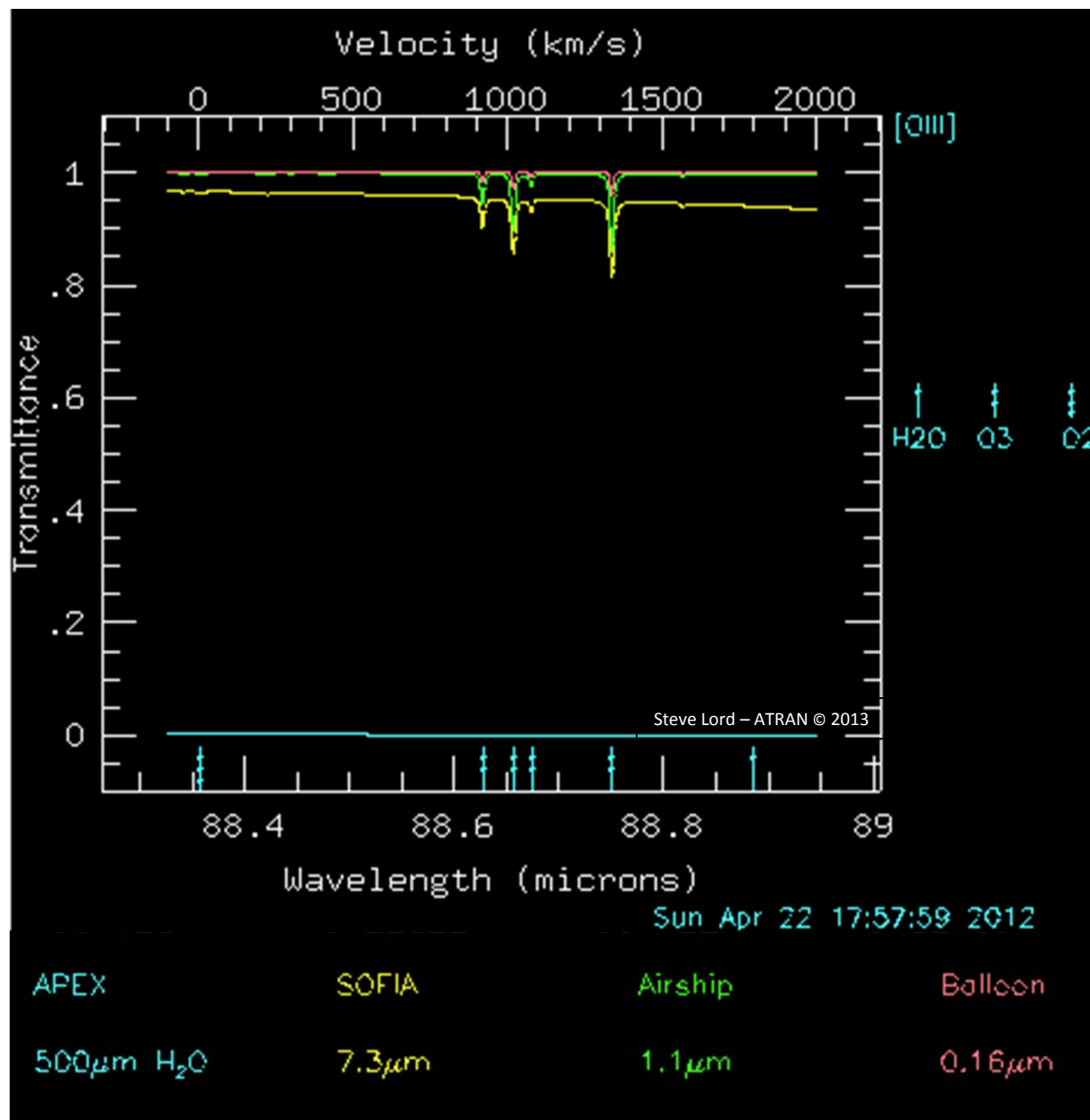
[FeIII] 23um



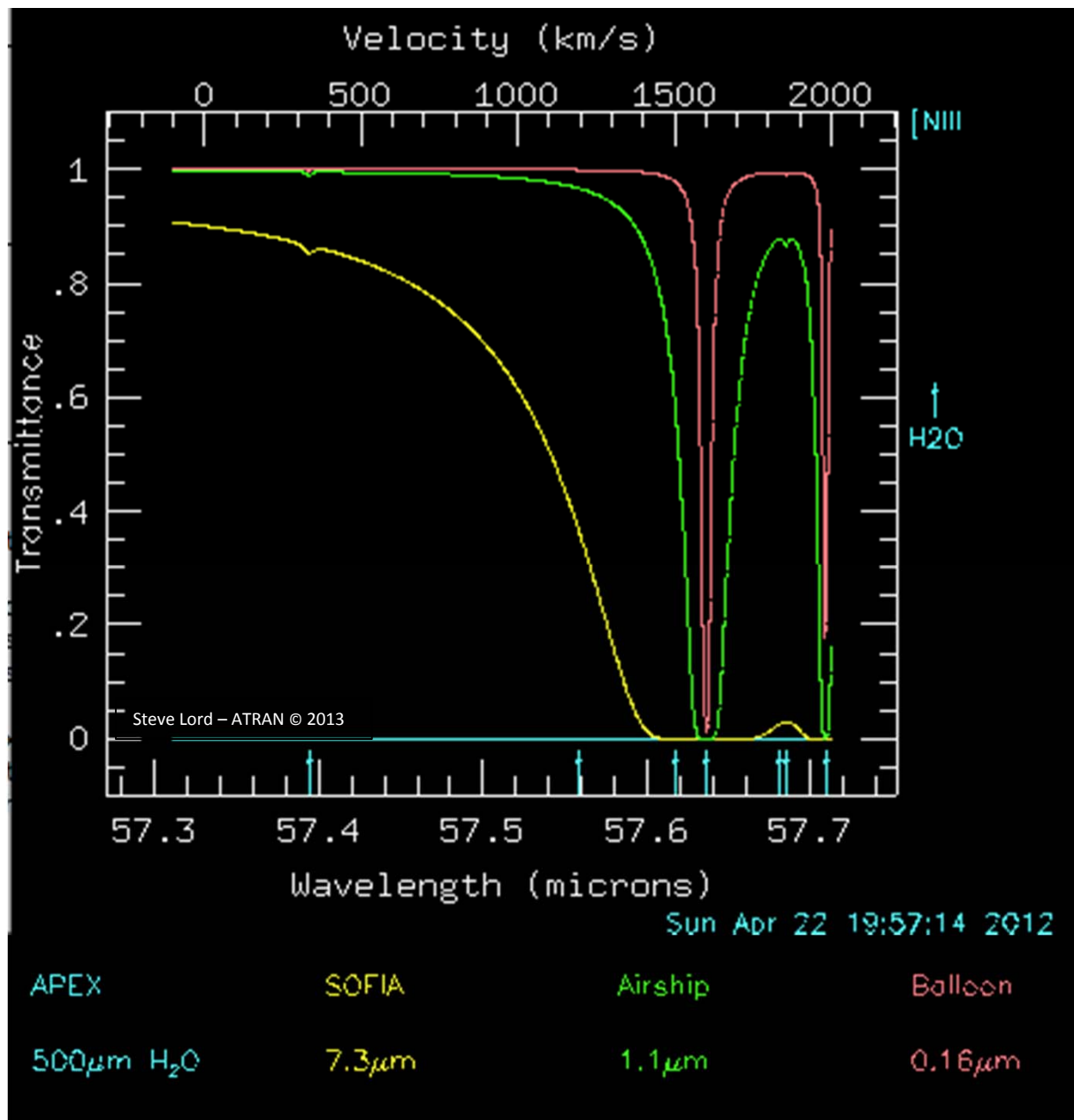
[NeIII] 36 μm

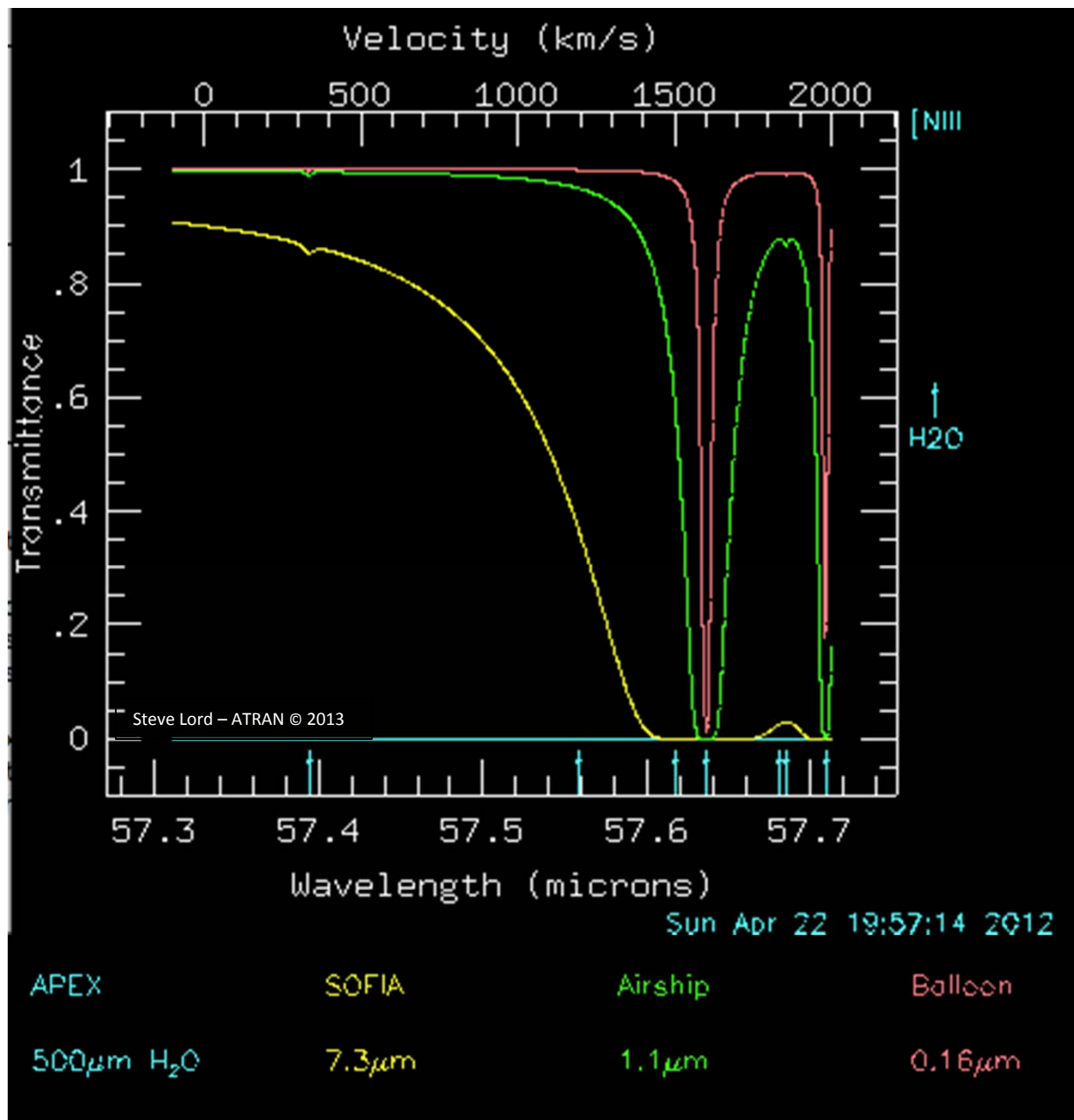
H₂O





[OIII] 88um



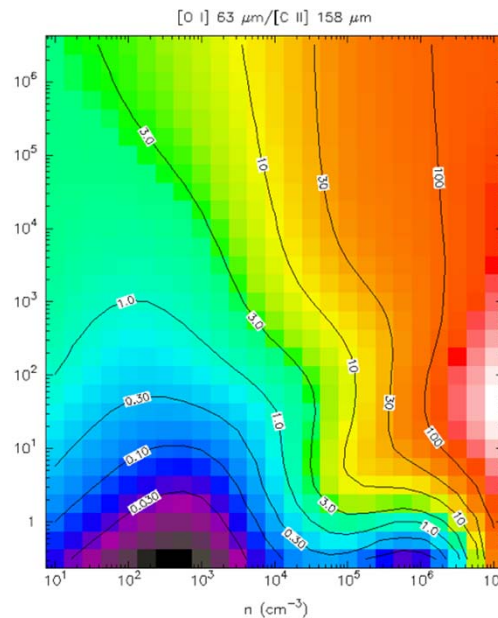
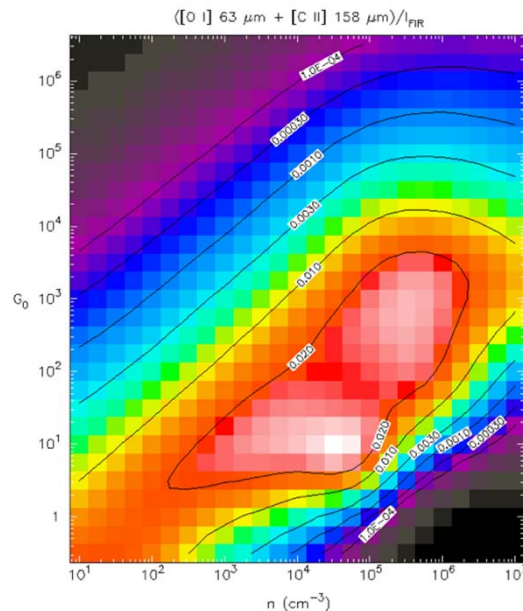


[NIII] 57 μ m

↑
H₂O

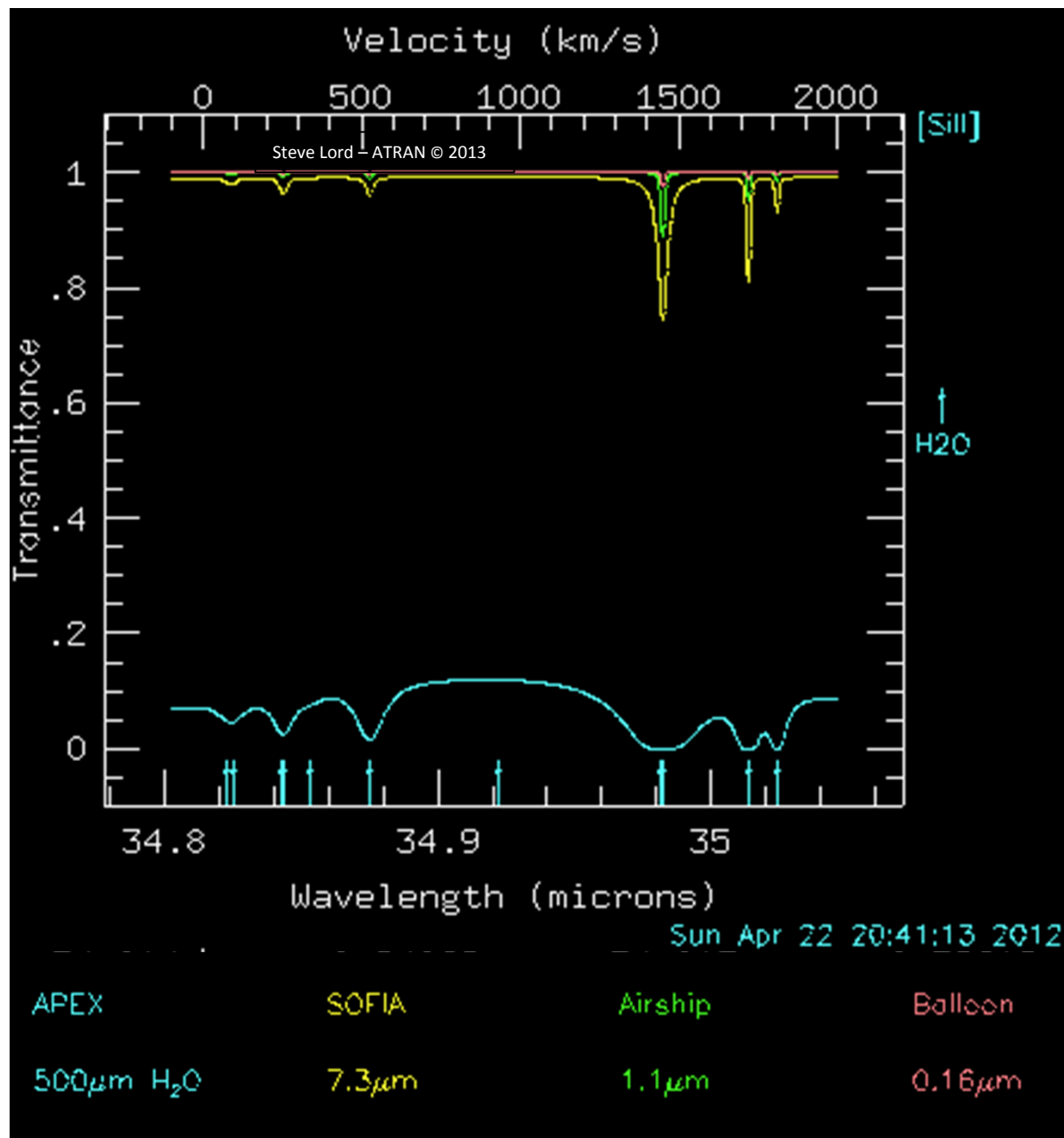
Photodissociation Region Diagnostics (PDRs)

- [SiII] 35 μ m (present in PDRs and HII regions)
- [OI] 63 μ m and 145 μ m
- [CII] 158 μ m (present in PDRs, HII regions, and the diffuse ISM)
– must be sorted out....

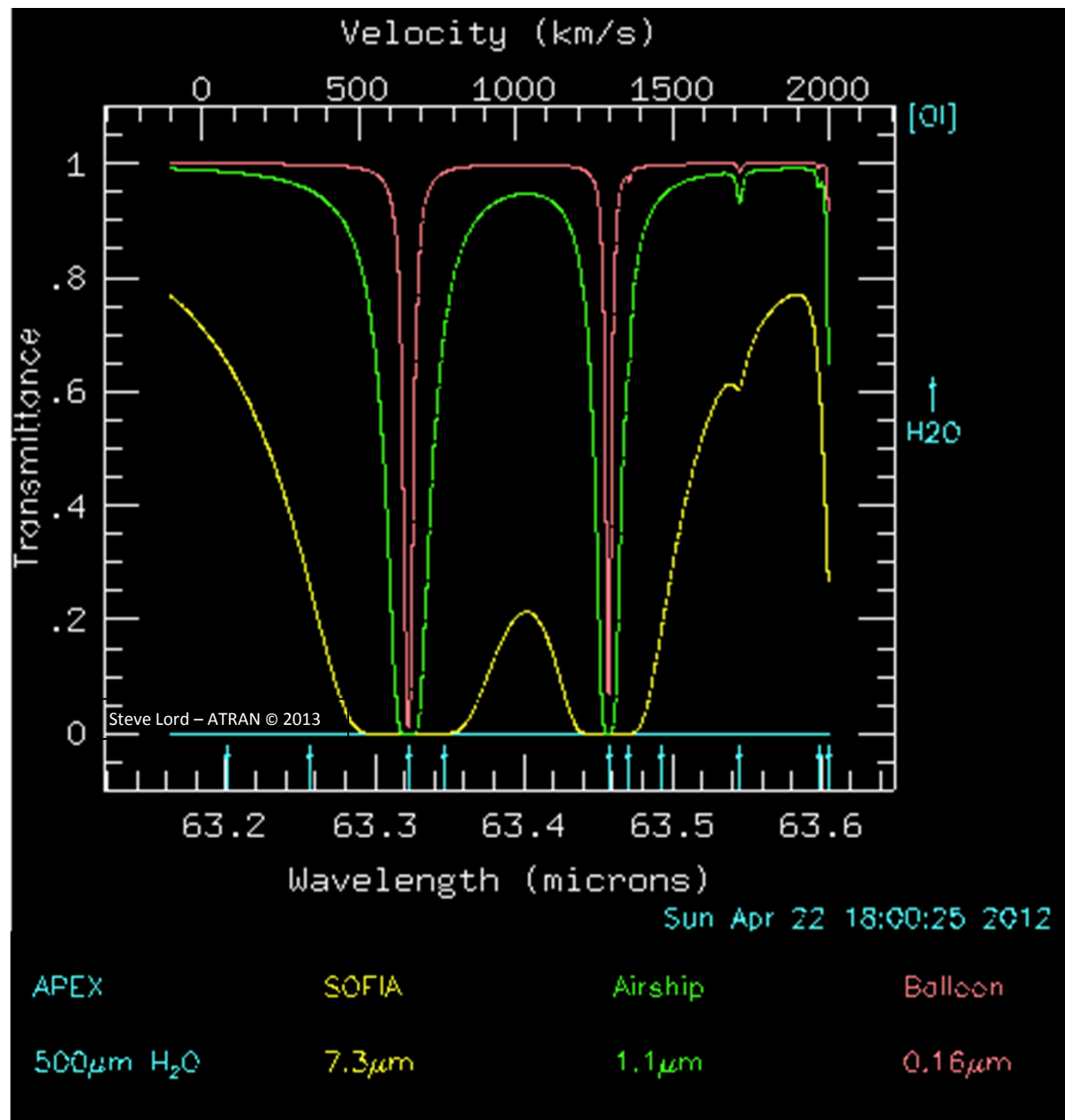


Give for the warm
neutral ISM:
 n (atomic density)
 G_0 (incident UV field)

Along with CO:
cloud sizes
cloud filling factors...

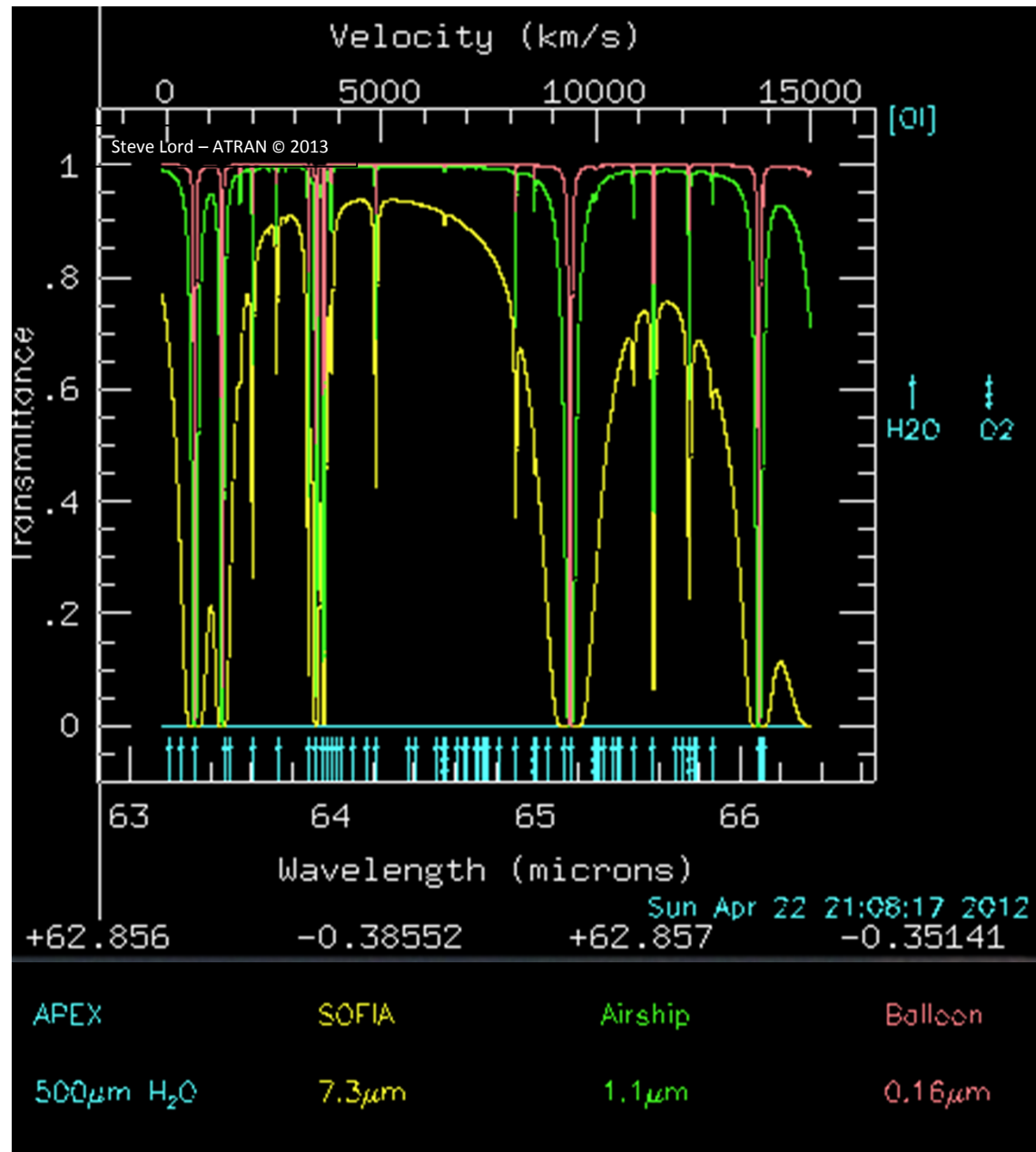


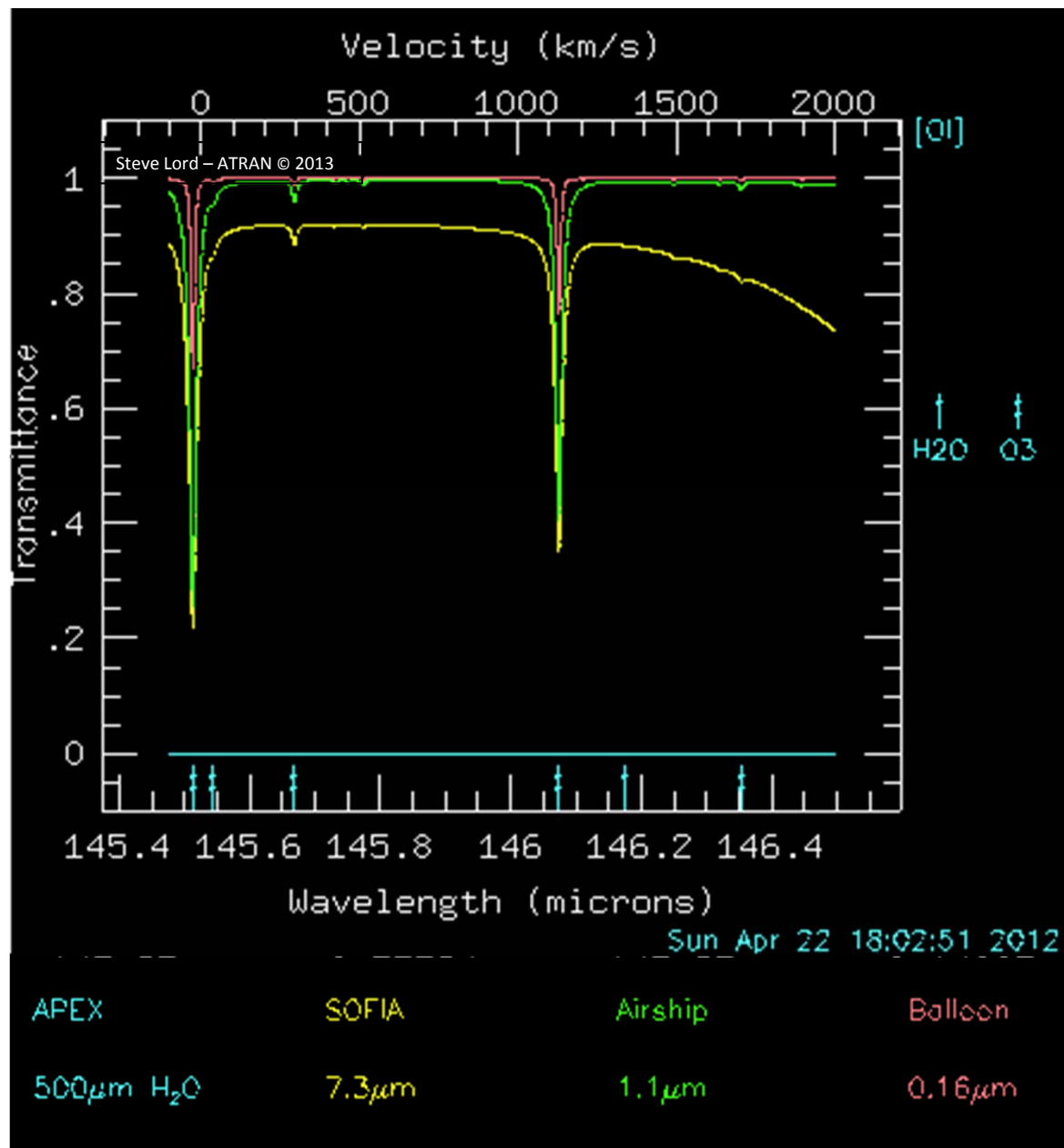
[SII] 35um



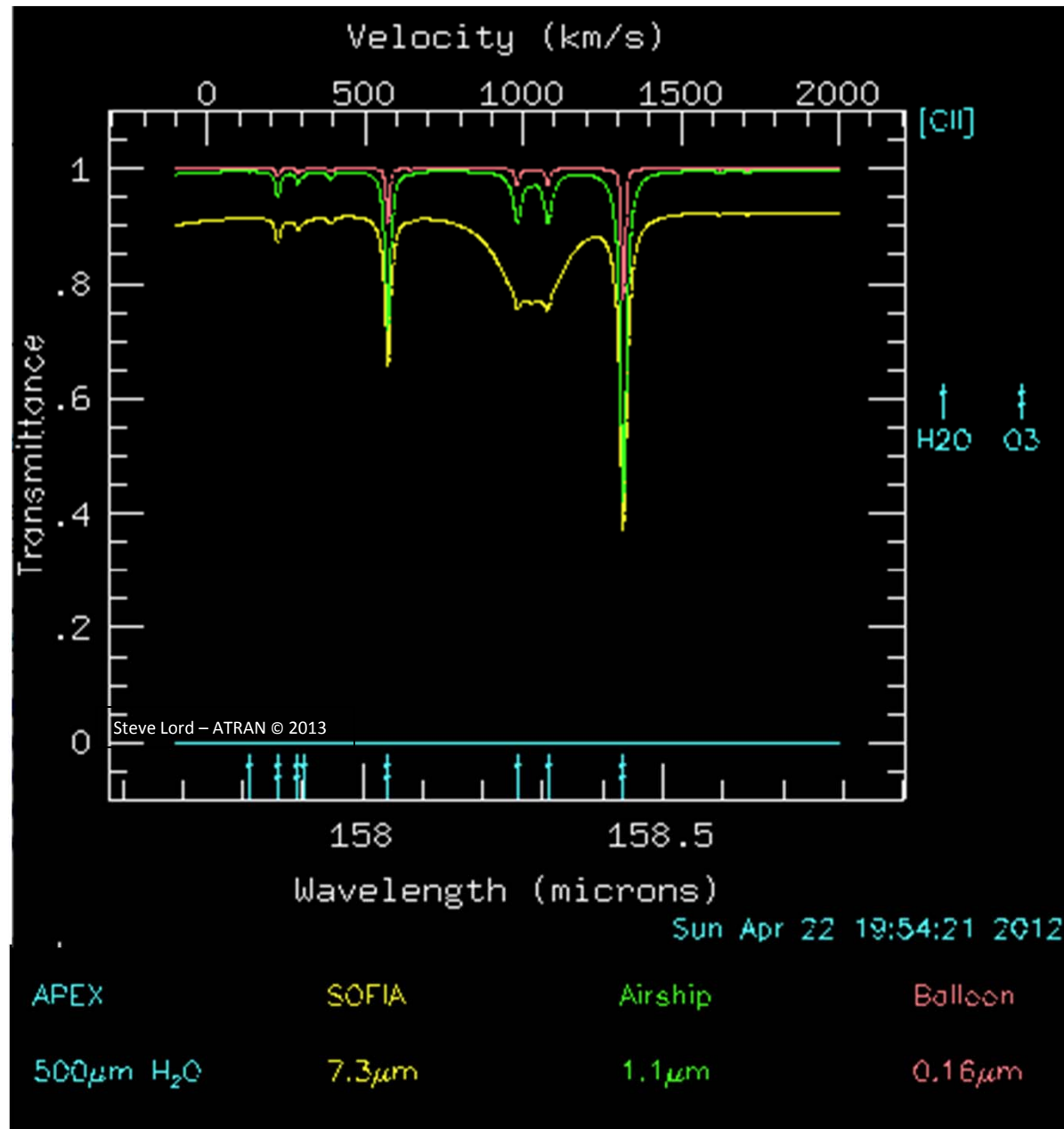
[OI] 63um

H₂O

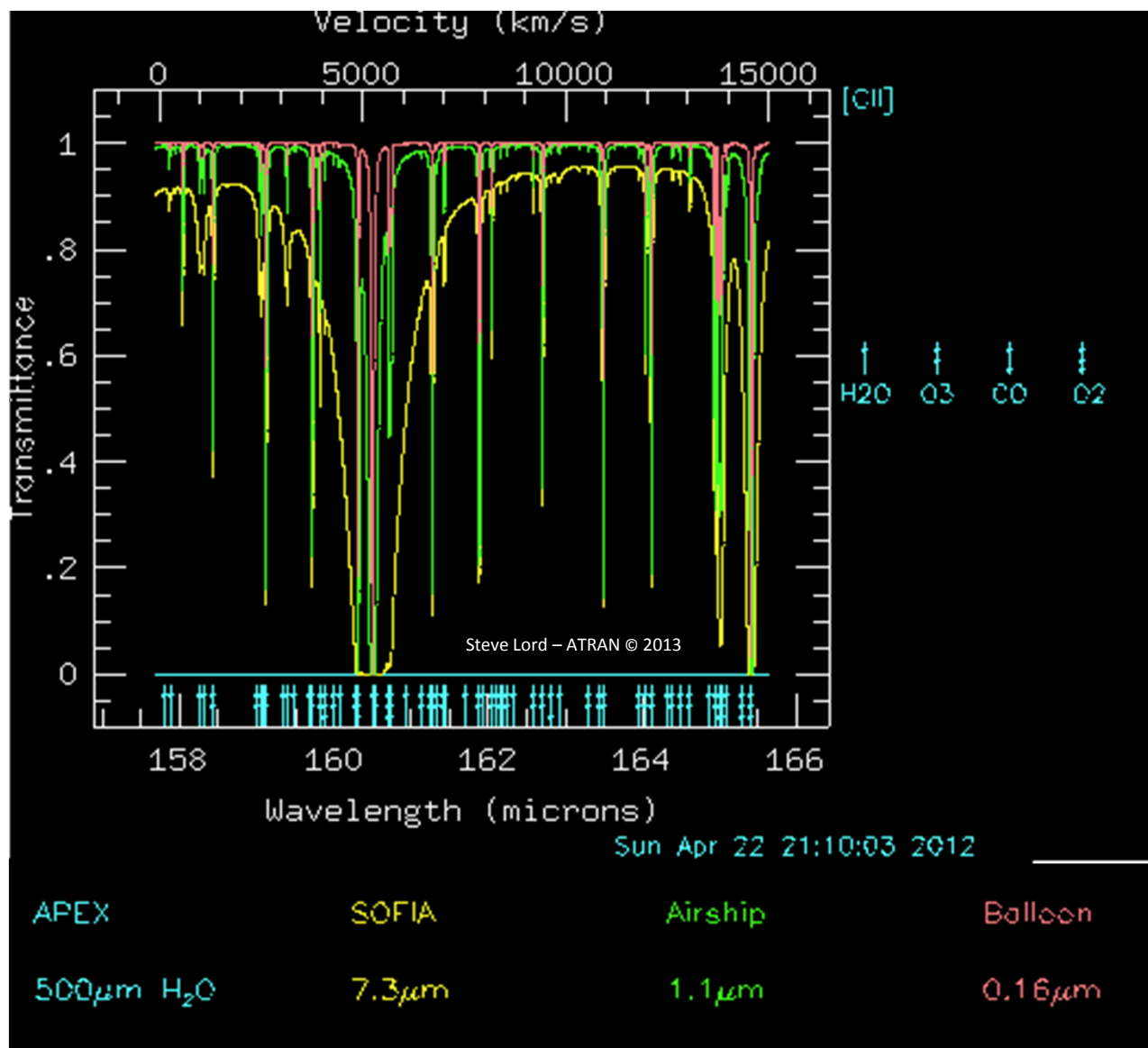




[OI] 145um



[ClII] 158um



[ClII] 158um
To:
v=15,000 km/s

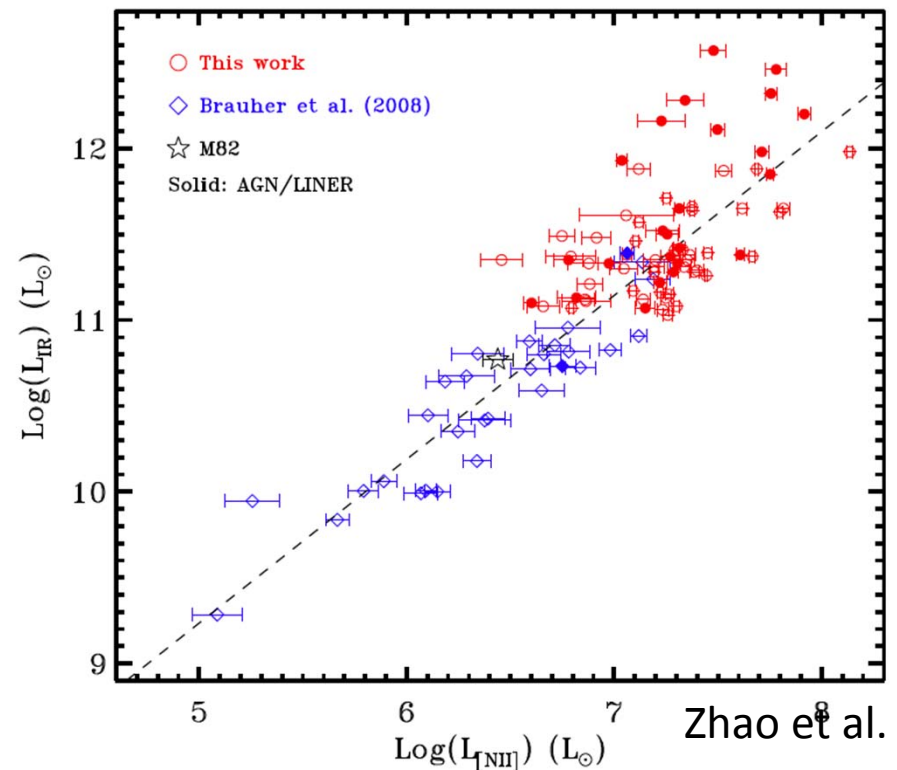
Diffuse ISM Diagnostics

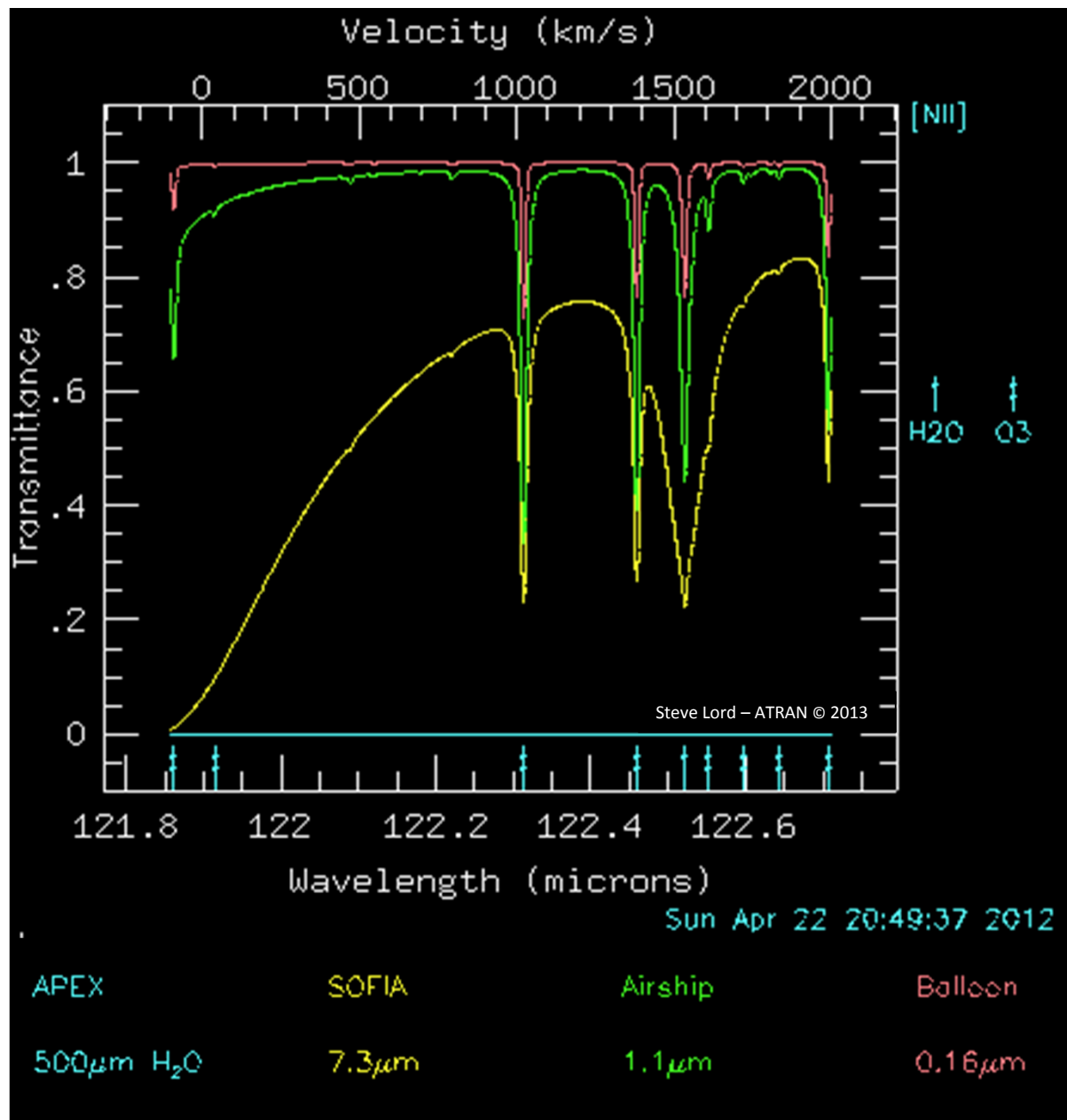
- [NII] 122 μ m and 205 μ m

Can measure of the density of diffuse ISM.

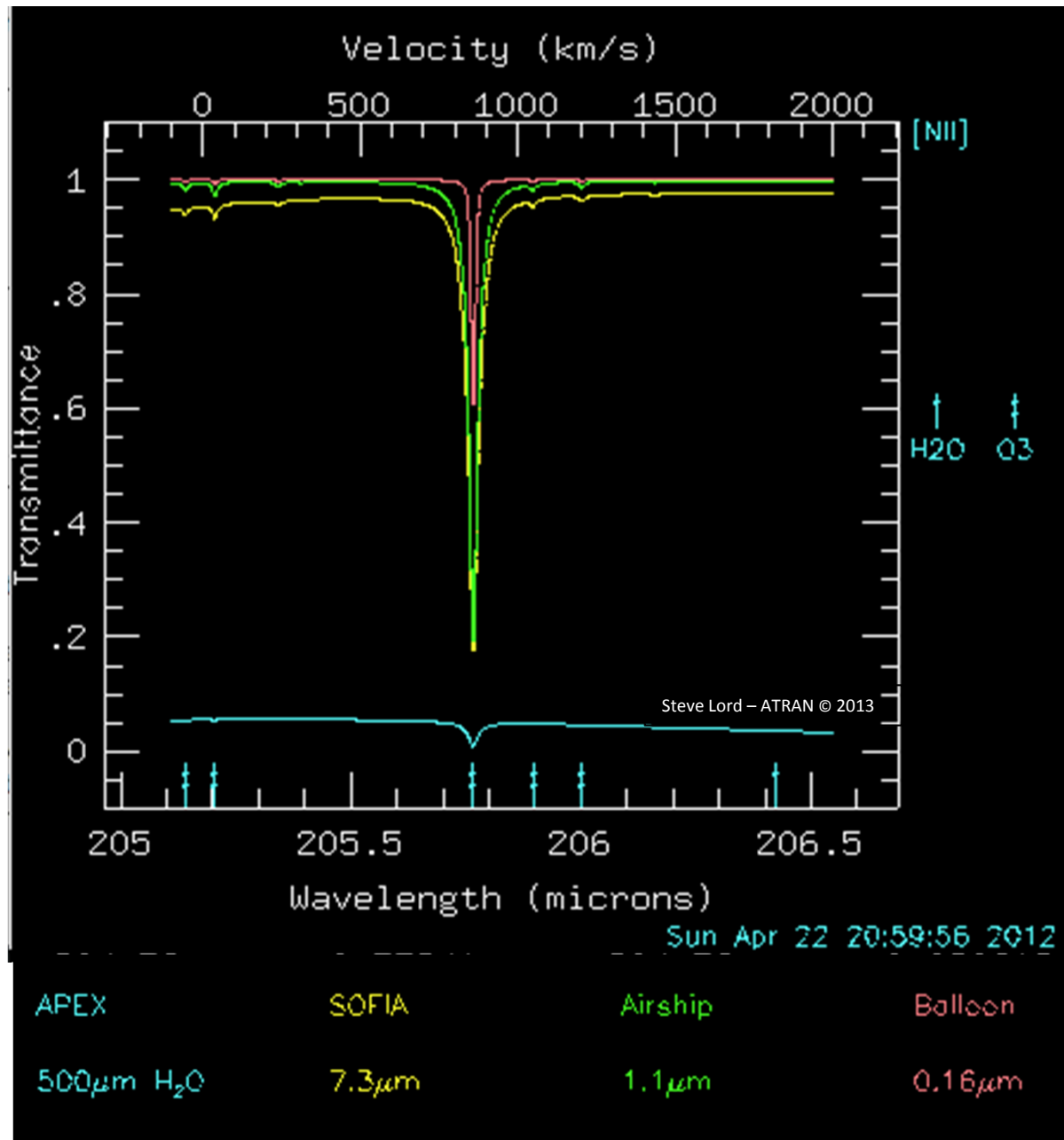
Can be used to subtract the diffuse [CII] component for PDR modeling.

Can serve as a proxy for L_{IR} and perhaps SFR (useful at high- z).





[NII] 122um



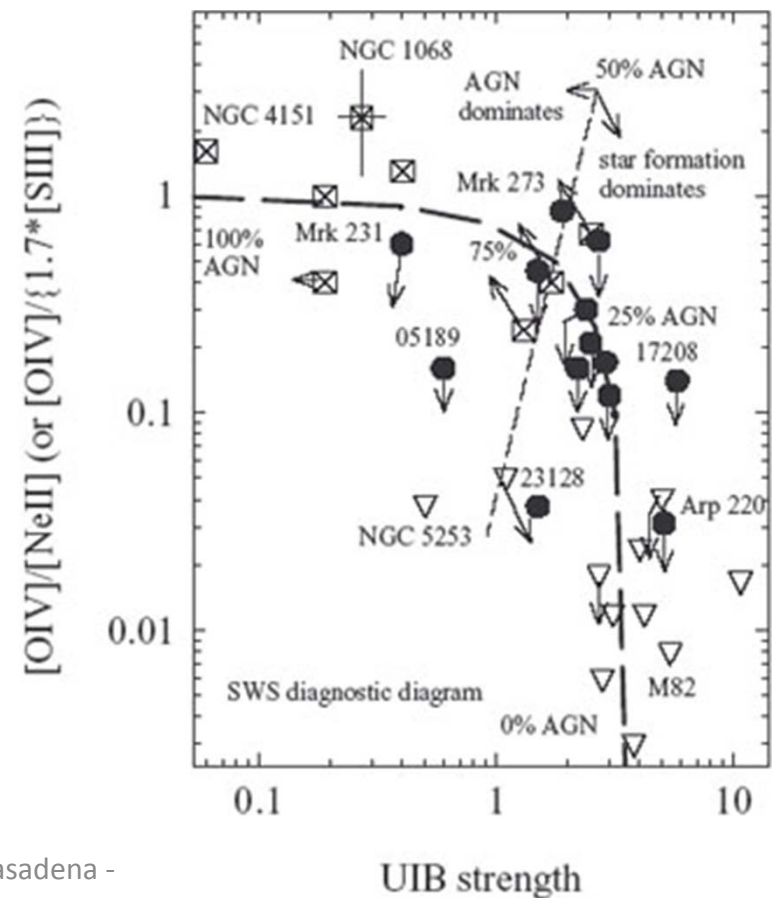
[NII] 205um

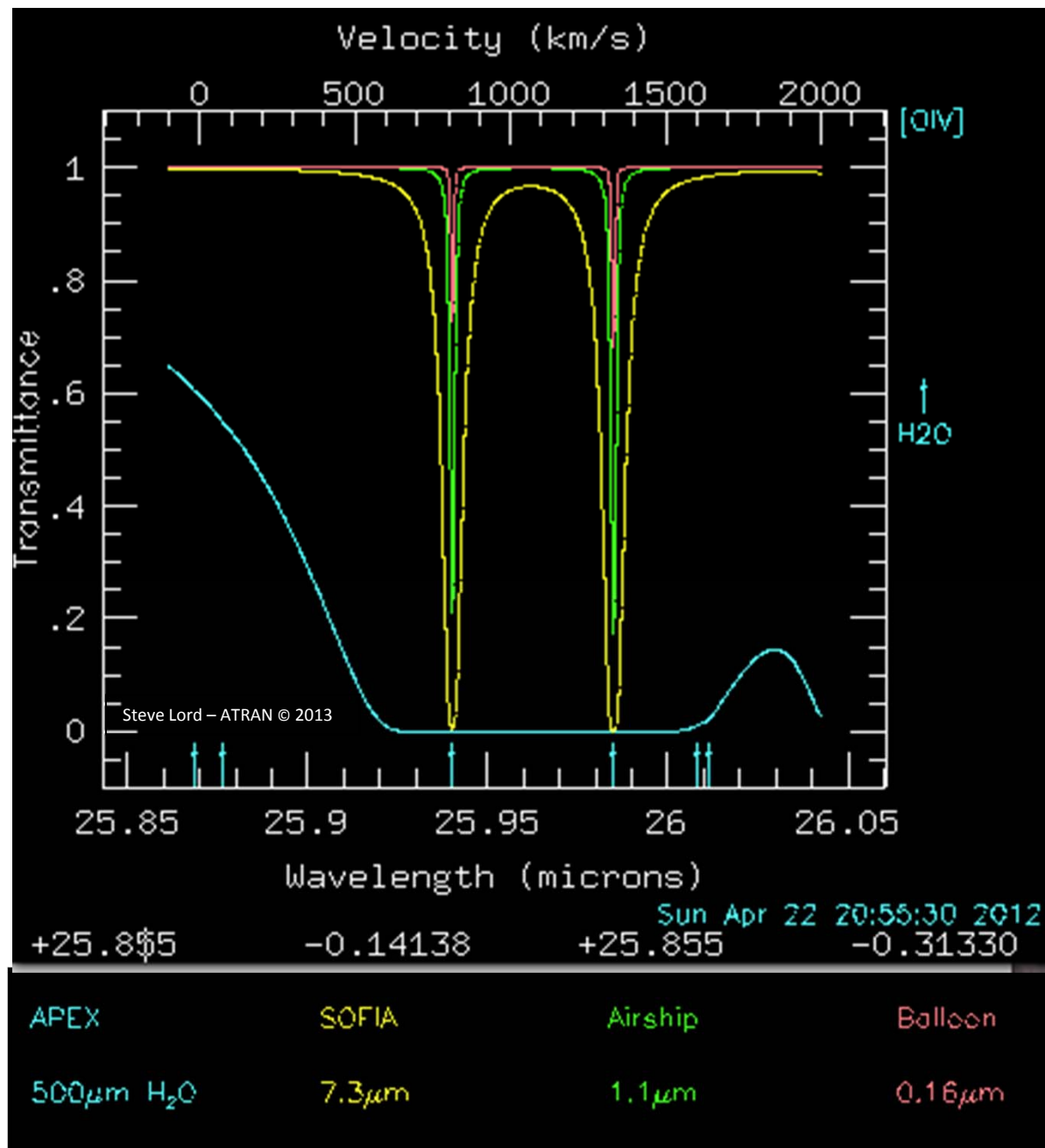
AGN Diagnostics

- [OIV] 26 μ m (along with [NeII] 13 μ m and PAHs)

Measure hardness of
radiation field

Distinguish AGNs and
Starbursts





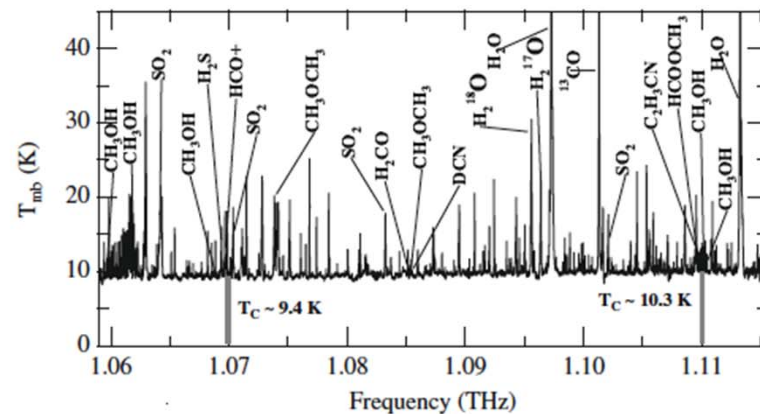
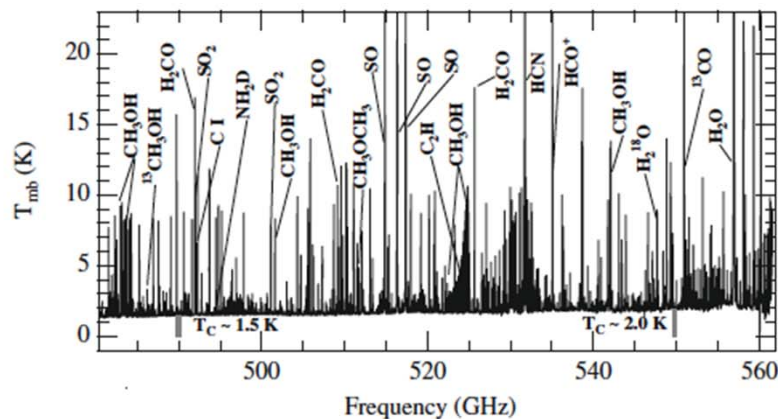
[OIV] 26um

Molecular Transitions

Example: HEXOS (Herschel Observations of Extraordinary Sources) –
T. Bergen, P.I. – including full coverage of Orion KL from 0.6-1.9 THz

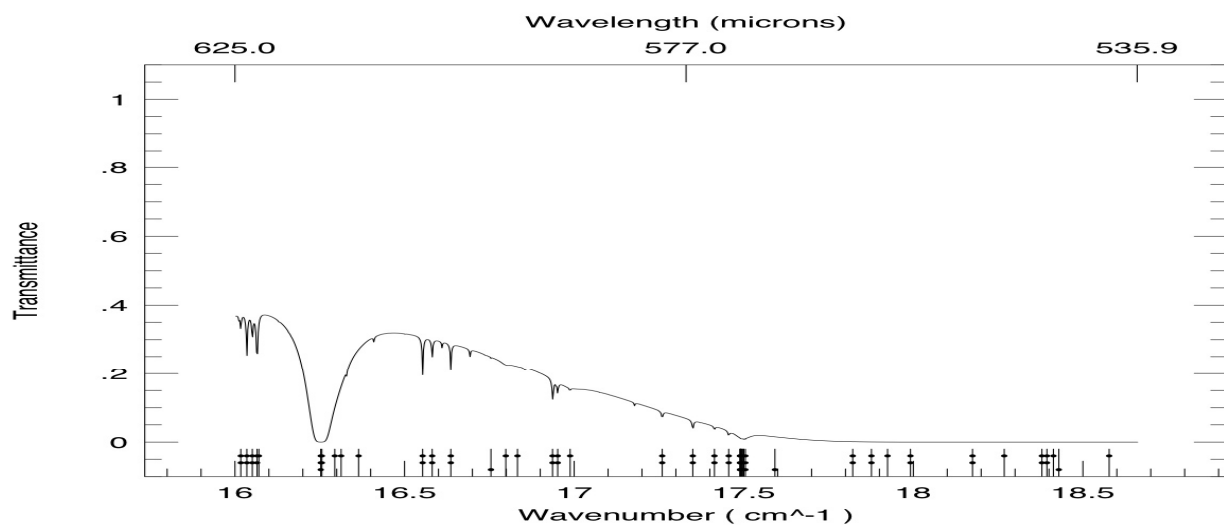
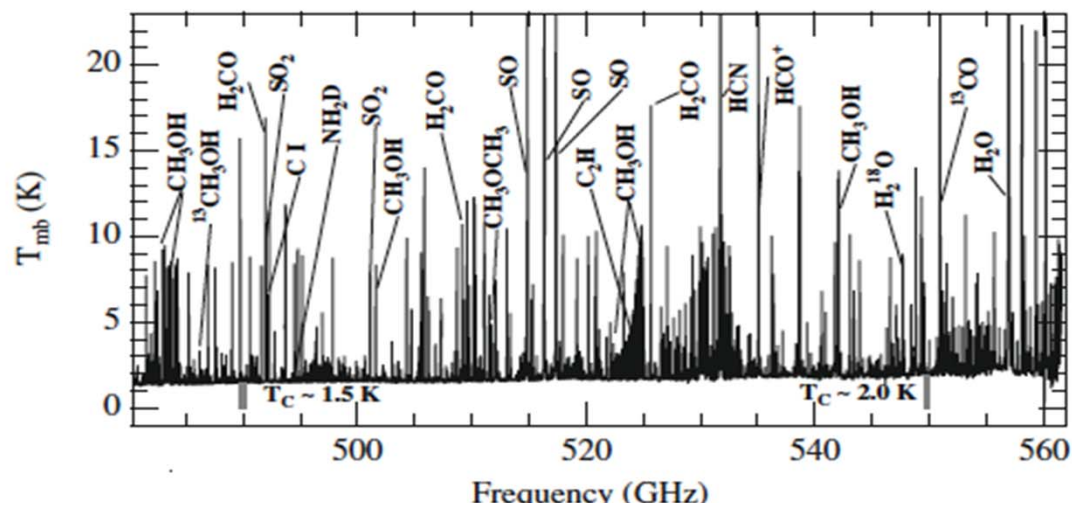
Resolves the dust continuum and dozens of lines in many hundreds of transitions (a line forest) allowing radiative transfer modeling of temperature, velocity, and density - specie by specie.

Look at these two samples – and look at the transmission through the Earth atmosphere.



These are Space Observations

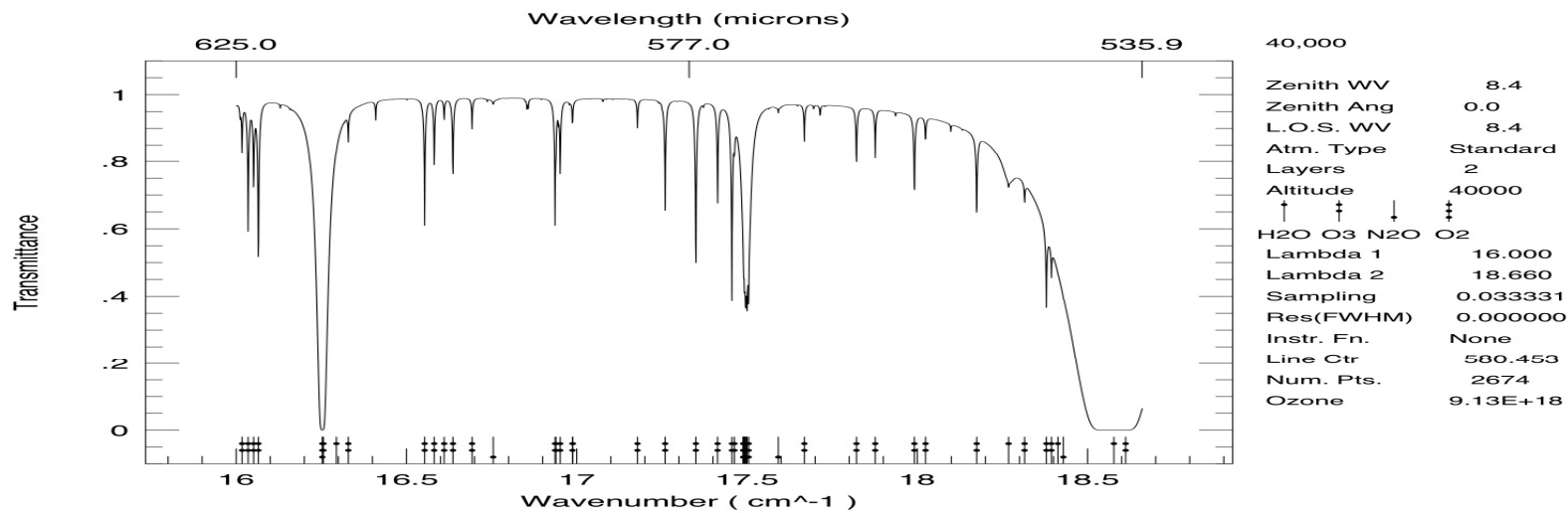
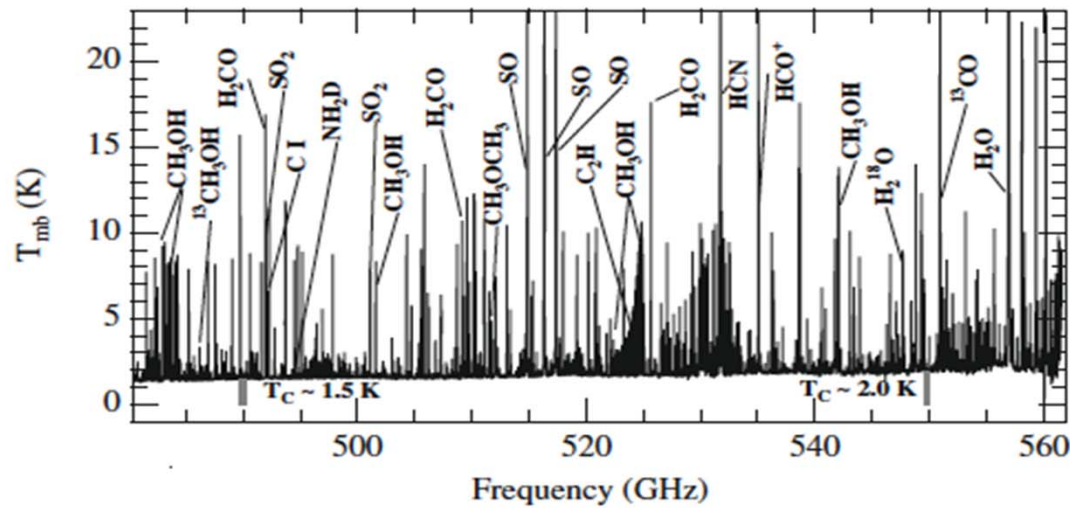
20,000' (577 μm)



Zenith WV	808.9
Zenith Ang	0.0
L.O.S. WV	808.9
Atm. Type	Standard
Layers	2
Altitude	20000
H2O	↑
O3	↑
N2O	↓
O2	↑
Lambda 1	16.000
Lambda 2	18.660
Sampling	0.079265
Res(FWHM)	0.000000
Instr. Fn.	None
Line Ctr	580.453
Num. Pts.	1125
Ozone	9.13E+18

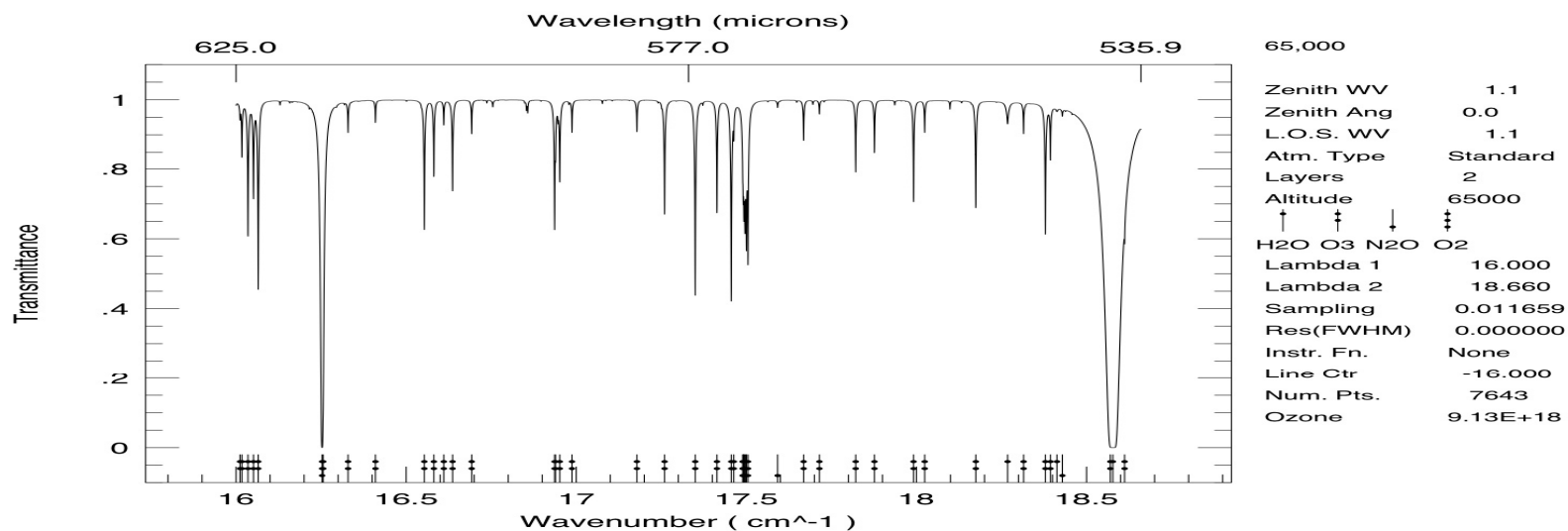
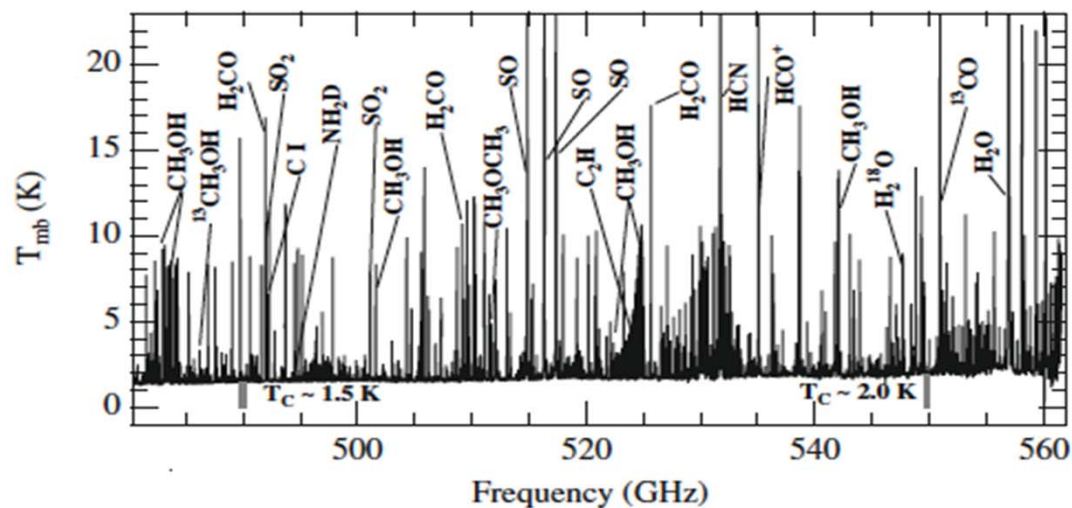
Tue Apr 30 04:46:17 2013

40,000' (577 μm)



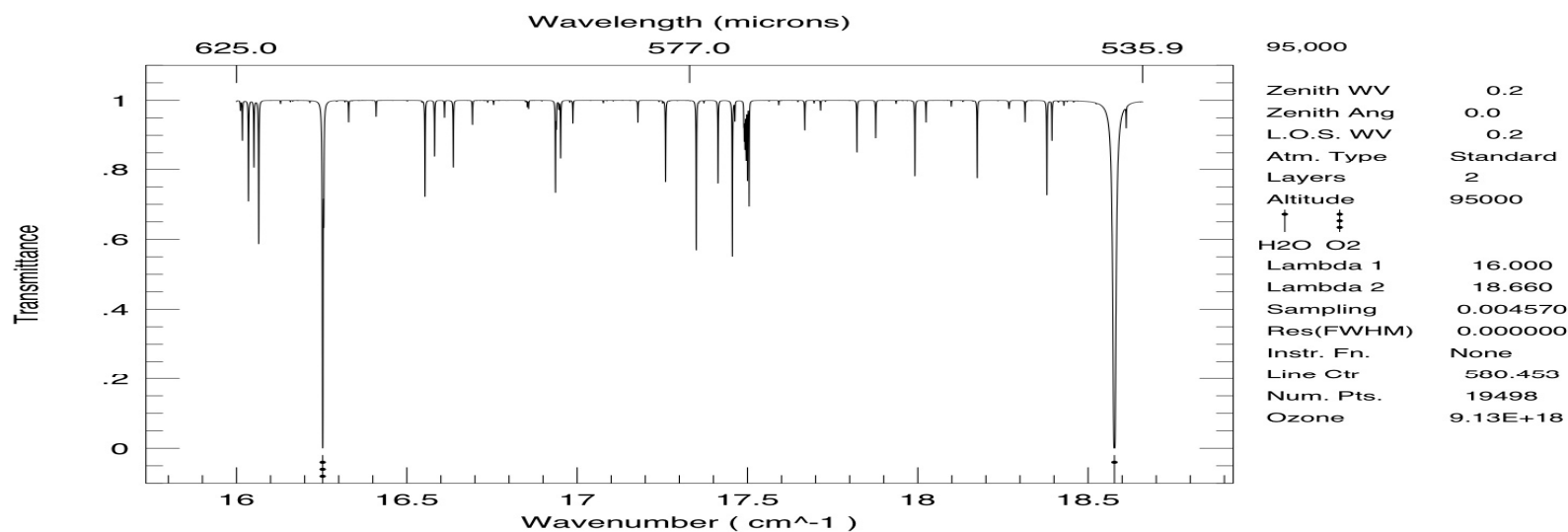
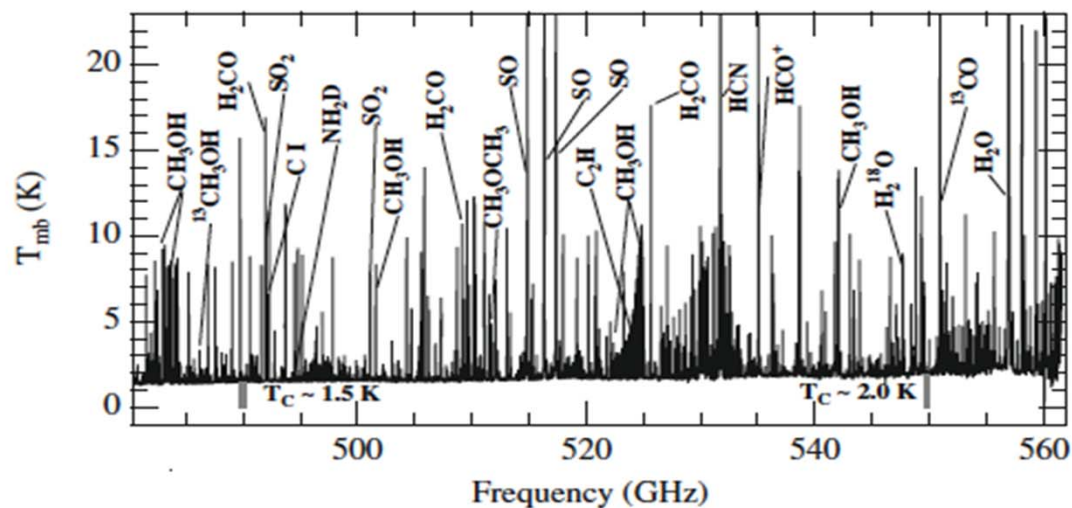
Tue Apr 30 04:48:28 2013

65,000' (577 μm)



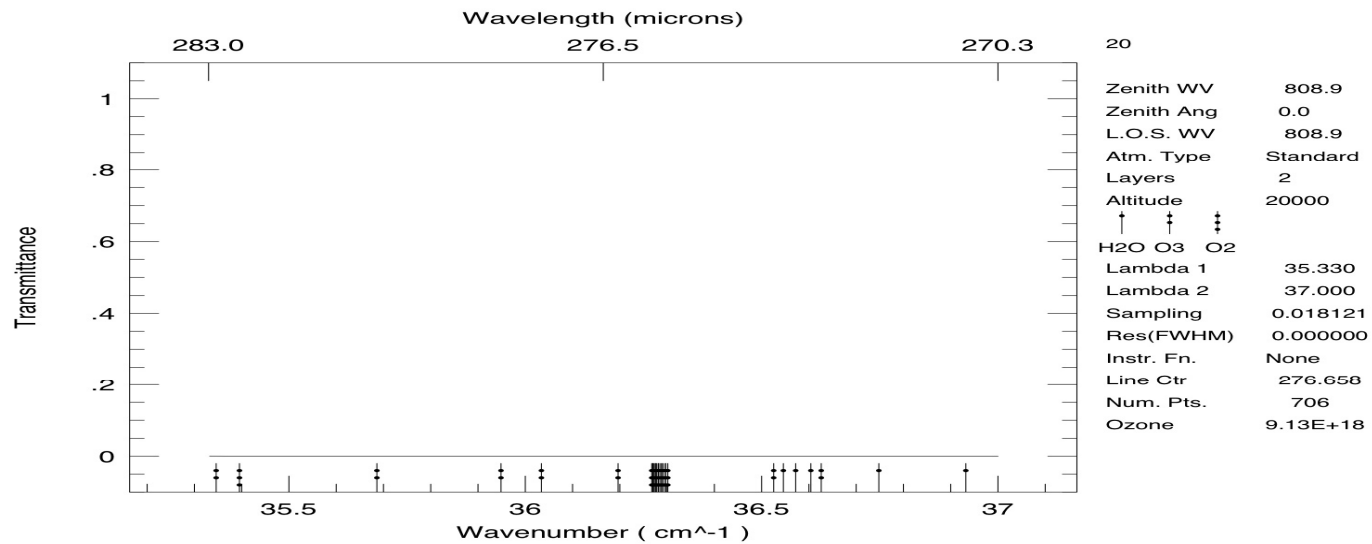
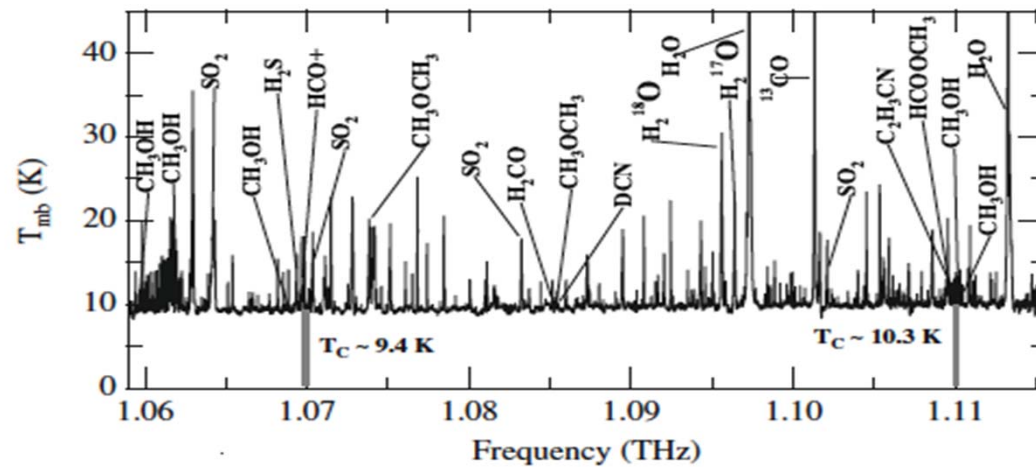
Tue Apr 30 04:51:01 2013

95,000' (577 μm)



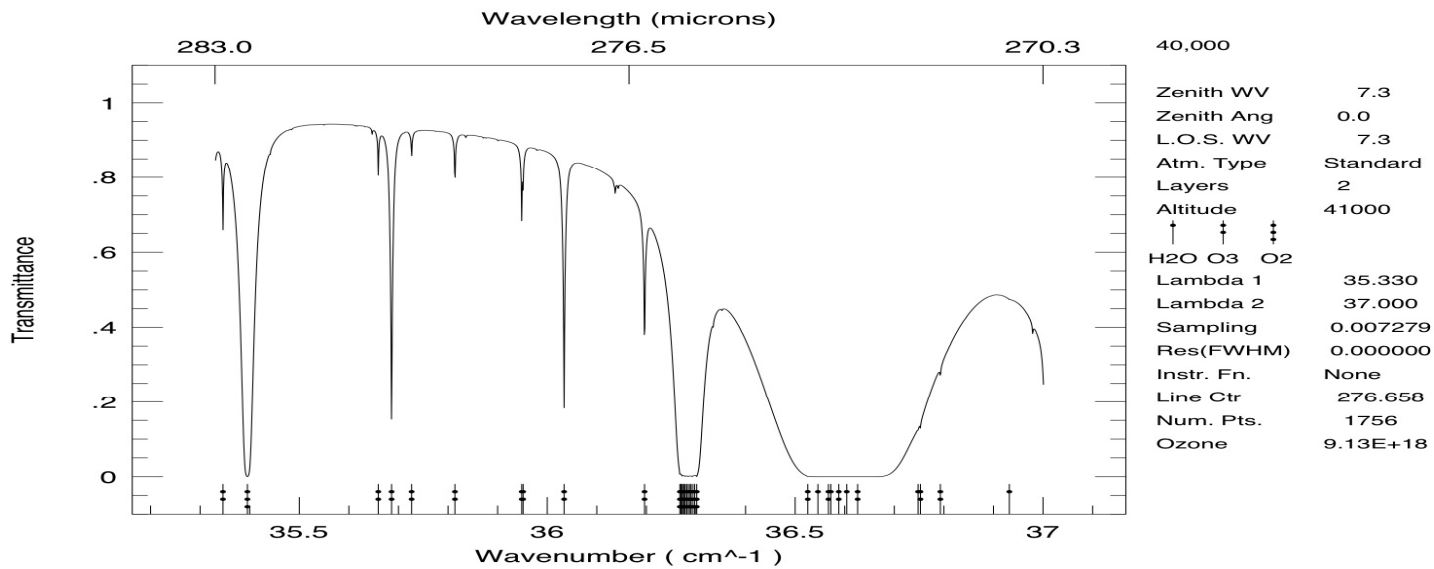
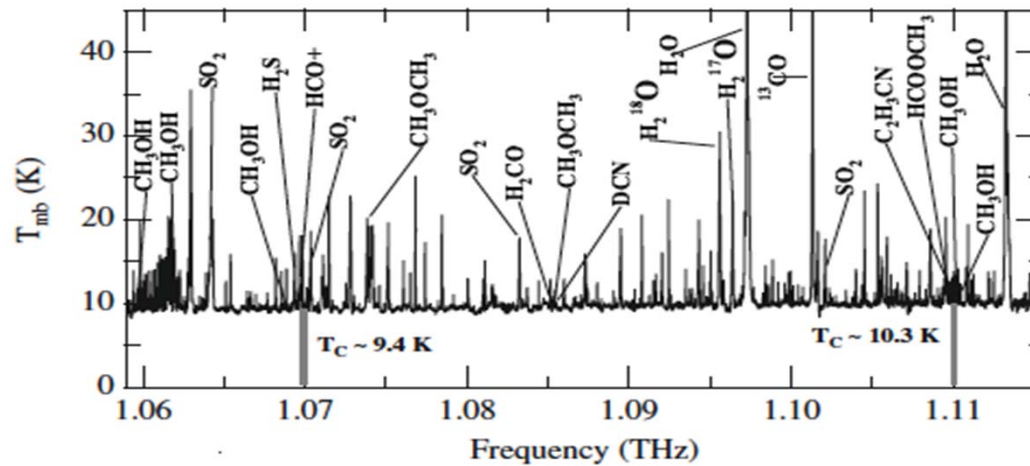
Tue Apr 30 04:53:08 2013

20,000' (276 μ m)



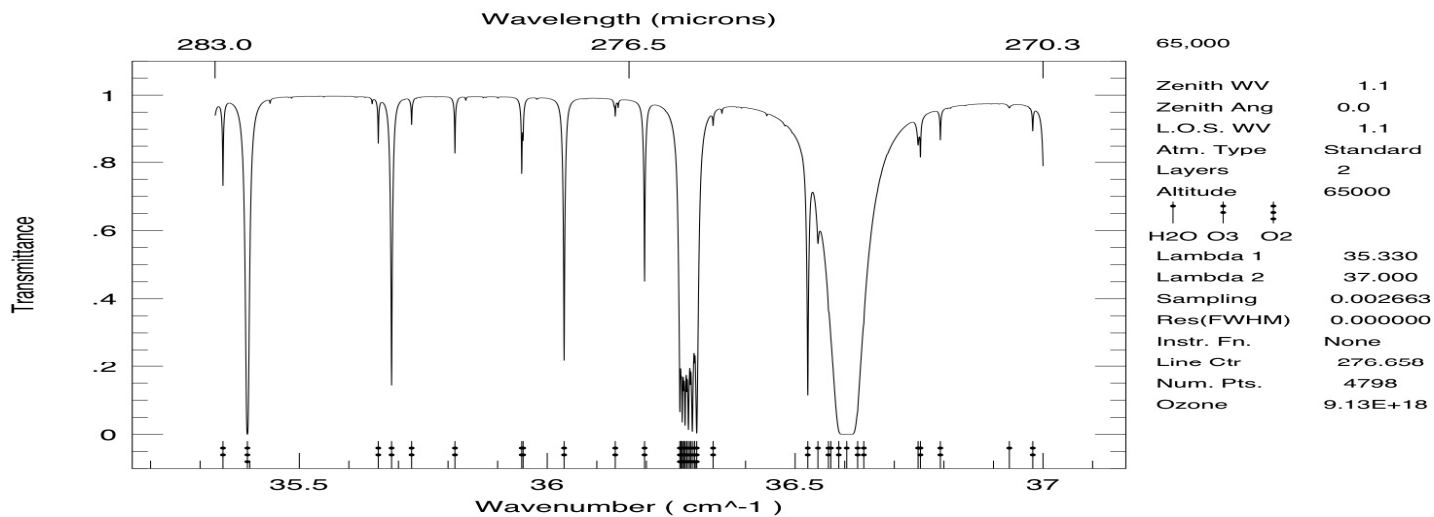
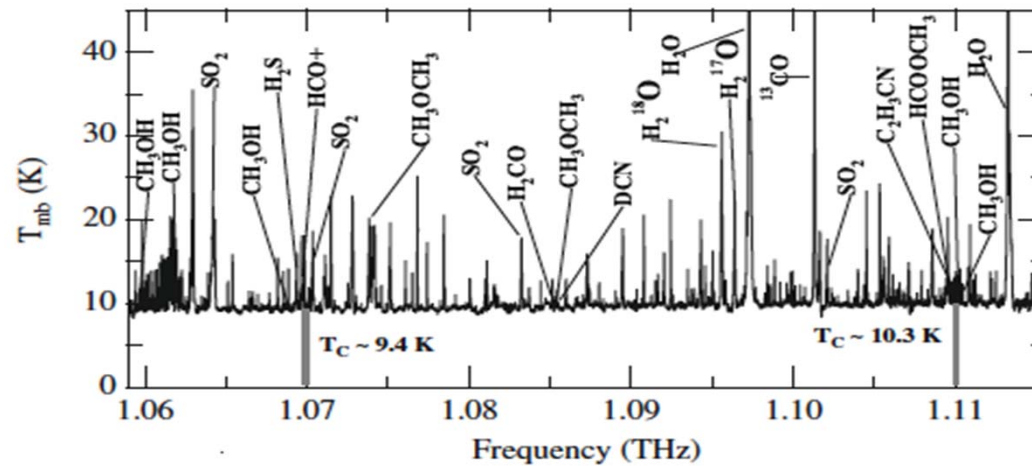
Tue Apr 30 04:25:51 2013

40,000' (276 μ m)



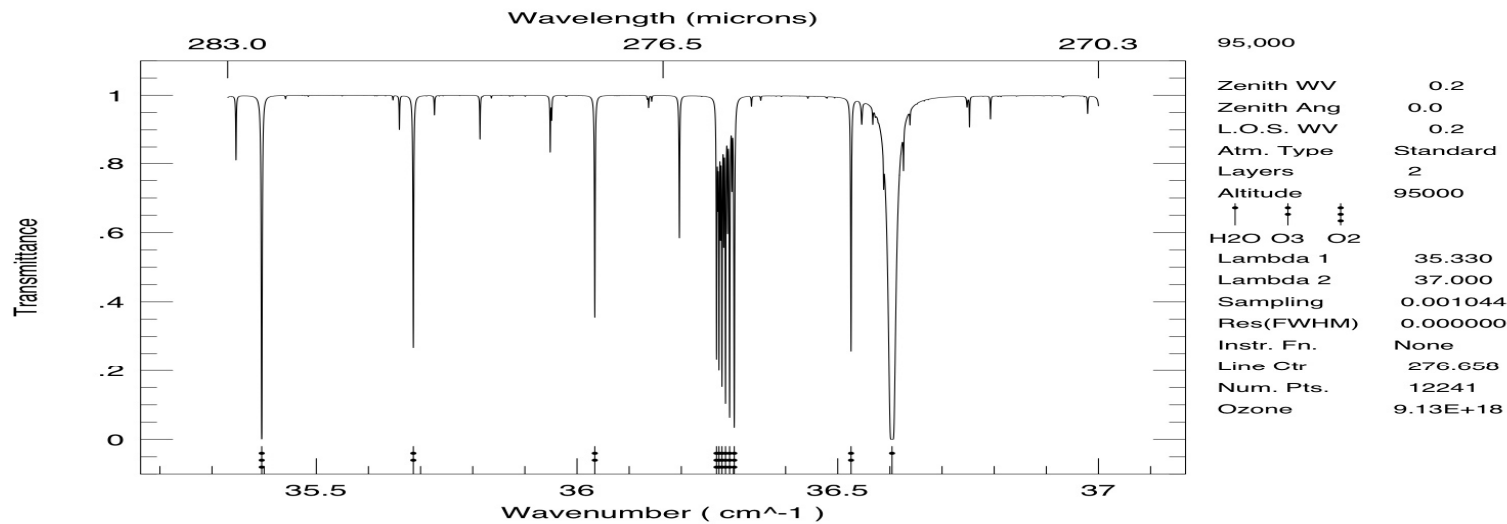
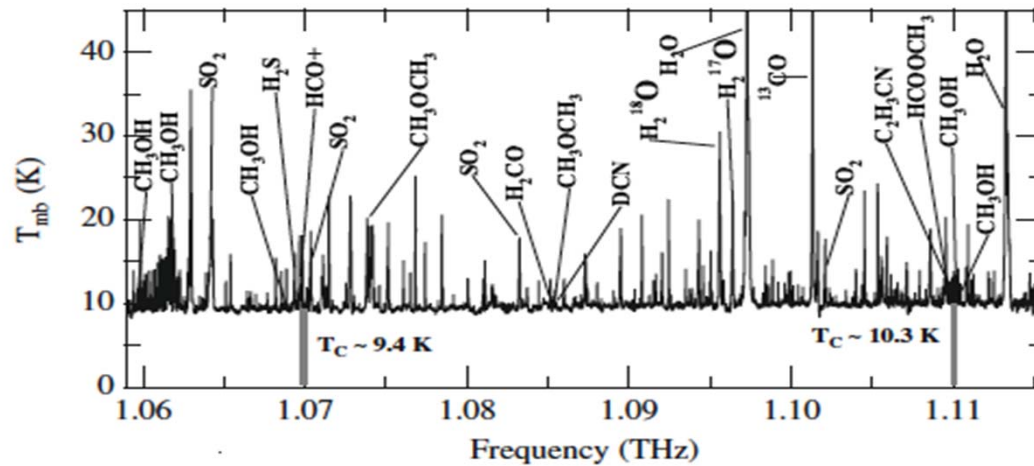
Tue Apr 30 04:30:01 2013

65,000' (276 μ m)



Tue Apr 30 04:32:28 2013

95,000' (276 μ m)



Tue Apr 30 04:37:07 2013

Hi-GAL

Herschel infrared Galactic Plane Survey

Molinari et al. 2010

Hi-GAL Team / ESA / PACS / SPIRE

PACS/SPIRE 70-160-350 μ m composite

Toward a Predictive Global Model of Galactic Star Formation

The *Hi-GAL* Team Institutes [PI: S. Molinari, INAF-IAPS Rome]

Italy: INAF-IAPS (Rome), Univ. Roma "Tor Vergata", Univ. Roma "La Sapienza", INAF-Oss. Arcetri, INAF-Oss. Catania, Univ. Salento. *USA*: Caltech-SSC-IPAC (Pasadena), Univ. Colorado (Boulder). *UK*: STFC-RAL (Chilton), Univ. Cardiff, Liverpool John Moores Univ., UCL (London), Univ. Hertfordshire, Univ. Leeds, Univ. Manchester, Univ. Exeter. *France*: CNRS-IRAP (Toulouse), LAM (Marseille), IAS (Orsay), CEA-Saclay (Saclay). *Canada*: Univ. Toronto, Univ. Calgary, Univ. Laval (Montreal). *Germany*: MPIfR (Bonn). *Japan*: Nagoya University. *China*: NAO-CAS (Beijing). *ESO-HQ* (Munich). *ESA-ESTEC* (Noordwijk). *ESA-ESAC* (Madrid).

Hi-GAL data processing is carried out at INAF-IAPS (Rome) thanks to support from Agenzia Spaziale Italiana under Contract I/038/08/0

Transmission of Herschel Continuum Filters (approx.)

	PACS	PACS	SPIRE
Filter	70um	160um	350um
65,000'	0.8	0.8	0.9
41,000'	0.6	0.6	0.7
20,000'	0	0	0.4